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The Cluster-EAGLE project: global properties of simulated clusters with resolved galaxies

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ABSTRACT

We introduce the Cluster-EAGLE (C-EAGLE) simulation project, a set of cosmological hydrodynamical zoom simulations of the formation of 30 galaxy clusters in the mass range of $10^{14} < M_{200}/M_{\odot} < 10^{15.4}$ that incorporates the *Hydrangea* sample of Bahé et al. (2017). The simulations adopt the state-of-the-art EAGLE galaxy formation model, with a gas particle mass of $1.8 \times 10^6 M_{\odot}$ and physical softening length of 0.7 kpc. In this paper, we introduce the sample and present the low-redshift global properties of the clusters. We calculate the X-ray properties in a manner consistent with observational techniques, demonstrating the bias and scatter introduced by using estimated masses. We find the total stellar content and black hole masses of the clusters to be in good agreement with the observed relations. However, the clusters are too gas rich, suggesting that the active galactic nucleus (AGN) feedback model is not efficient enough at expelling gas from the high-redshift progenitors of the clusters. The X-ray properties, such as the spectroscopic temperature and the soft-band luminosity, and the Sunyaev–Zel’dovich properties are in reasonable agreement with the observed relations. However, the clusters have too high central temperatures and larger-than-observed entropy cores, which is likely driven by the AGN feedback after the cluster core has formed. The total metal content and its distribution throughout the intracluster medium are a good match to the observations.

Key words: hydrodynamics – methods: numerical – galaxies: clusters: general – galaxies: clusters: intracluster medium – X-rays: galaxies: clusters.

1 INTRODUCTION

The observable properties of galaxy clusters emerge from the complex interplay of astrophysical processes and gravity acting on hierarchically increasing scales. Cluster formation is a process that has an enormous dynamic range, as clusters collapse from fluctuations with a comoving scale length of tens of Mpc, but have observable properties that are shaped by highly energetic astrophysical

processes acting on subparsec scales (see Voit 2005; Allen, Evrard & Mantz 2011; Kravtsov & Borgani 2012, for recent reviews). The interaction of processes acting on very different scales makes cluster formation a highly non-linear process. However, the combination of scales and processes make galaxy clusters a unique environment where we can observe not only the material that participated in galaxy formation but also the material that did not. Therefore, clusters allow the simultaneous study of fundamental cosmological parameters, gravity, hydrodynamical effects, chemical element synthesis and the interaction of relativistic jets with the cluster environment.

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Although significant progress can be made with semi-analytic prescriptions (Bower, McCarthy & Benson 2008; Somerville et al. 2008; Guo et al. 2011; Bower, Benson & Crain 2012), hydrodynamical simulations are the only method that can capture the effects of physical processes during cluster formation and predict the resulting observable consequences self-consistently. Although unable to capture the full dynamic range due to limited computational resources, there has been significant progress in the modelling of cluster formation and the physical processes that occur below the resolution scale of the simulation, so-called subgrid models. The formation of the baryonic component of clusters has been well studied (e.g. Eke, Navarro & Frenk 1998; Kay et al. 2004; Crain et al. 2007; Nagai, Kravtsov & Vikhlinin 2007b; Sijacki et al. 2007; Dubois et al. 2010; Short et al. 2010; Young et al. 2011; Battaglia et al. 2012; Gupta et al. 2016), including the importance of feedback from active galactic nuclei (AGNs) and its effect on the baryonic content of clusters (e.g. Puchwein, Sijacki & Springel 2008; Fabjan et al. 2010; McCarthy et al. 2010; Martizzi et al. 2016). These developments have led to several independent groups simulating samples of clusters that are to varying degrees realistic (Le Brun et al. 2014; Pike et al. 2014; Planelles et al. 2014; Rasia et al. 2015; Hahn et al. 2017), i.e. their observable properties, such as X-ray luminosity and spectroscopic temperature, are a good match to those of observed clusters. The understanding that subgrid models should be calibrated against carefully selected observable relations has resulted in simulations that simultaneously reproduce a host of stellar, gas and halo properties (McCarthy et al. 2017), even to high redshift (Barnes et al. 2017). One limitation of previous cluster formation simulation work is that it only achieved a modest resolution, typically with a gas particle mass of $m_{\text{gas}} \sim 10^9 M_{\odot}$ [for smoothed particle hydrodynamic (SPH) simulations] and a spatial resolution of ~ 5 kpc. This limits the ability to resolve the structures in the intracluster medium (ICM), to examine the interactions between energetic astrophysical processes and the ICM, to capture the formation and evolution of the cluster galaxy population, and to resolve the growth histories of the black holes (BHs).

At the same time, there has been significant progress in the theoretical modelling of galaxy formation in representative volumes. Improved resolution and the development and calibration of efficient subgrid prescriptions for feedback processes have led to a step change in the realism of galaxy formation models (Vogelsberger et al. 2014; Crain et al. 2015; Schaye et al. 2015; Davé, Thompson & Hopkins 2016; Tremmel et al. 2017). For example, the EAGLE simulation suite (Schaye et al. 2015; Crain et al. 2015) was calibrated against the observed galaxy stellar mass function, the field galaxy size–mass relation and the BH mass–stellar mass relation at low redshift. Following this, the model then yields broad agreement with, among other things, the observed evolution of galaxy star formation rates (Furlong et al. 2015), the evolution of galaxy sizes (Furlong et al. 2017), their molecular and atomic hydrogen content (Lagos et al. 2015; Bahé et al. 2016; Crain et al. 2017), their observed colour distribution (Trayford et al. 2015), and the growth of BHs and their link to the star formation and growth of galaxies (Rosas-Guevara et al. 2016; Bower et al. 2017; McAlpine et al. 2017). However, the resolution required and complexity of the subgrid models make these simulations computationally expensive, limiting their volume to periodic cubes with a side length of ~ 100 Mpc or less. Although a volume of this size will contain many galaxy groups ($M_{200} = 10^{13}–10^{14} M_{\odot}^1$), rich galaxy clusters (M_{200}

$\geq 10^{15} M_{\odot}$) are very rare objects and a volume of this size is highly unlikely to contain even one. Therefore, it is difficult to assess the ability of these calibrated models to produce realistic large-scale structures, such as galaxy clusters, and to test whether they correctly capture galaxy formation in the full range of environments.

Motivated by the limitations of existing cluster and galaxy formation simulations, we introduce the Virgo consortium’s Cluster-EAGLE (C-EAGLE) project. The project consists of zoom simulations of the formation of 30 galaxy clusters that are evenly spaced in the mass range of $10^{14}–10^{15.4} M_{\odot}$, probing environments that are not present in the original periodic EAGLE volumes presented by Schaye et al. (2015), henceforth S15, and Crain et al. (2015). They are performed with the EAGLE galaxy formation model (AGNt9 calibration) and adopt the same mass resolution ($m_{\text{gas}} = 1.81 \times 10^6 M_{\odot}$) and physical spatial resolution ($\epsilon = 0.7$ kpc) as the largest periodic volume of the EAGLE suite (Ref-L100N1504). The resolution of the simulations allows us to resolve the formation of cluster galaxies and their co-evolution with the ICM, the interactions between the cluster galaxies and the ICM, the formation of structures within the ICM, and how energetic astrophysical processes shape the ICM. As the hot halo typically extends to several virial radii, the zoom regions extend to at least five times the virial radius of each object to include the large-scale structure around them. The *Hydrangea* sample (Bahé et al. 2017), designed to study the evolution of galaxies as their environment transitions from isolated field to dense cluster, extends the zoom region to 10 virial radii for 24 of 30 C-EAGLE clusters.

In this paper, we present the global properties and hot gas profiles of the simulated clusters at low redshift and compare to observations in order to examine the ability of a model calibrated for galaxy formation to produce realistic galaxy clusters. In a companion paper (Bahé et al. 2017), the properties of the cluster galaxy population are presented and the *Hydrangea* sample is used to study the impact of the cluster environment on the galaxy stellar mass function. The predicted galaxy luminosity functions of the clusters will be presented in Dalla Vecchia et al. (in preparation), including results for higher-resolution runs of a subset of the clusters. The rest of this paper is structured as follows. In Section 2, we present the sample selection, a brief overview of the EAGLE model and the method adopted for computing global properties in a manner consistent with observational techniques, which enables a fairer comparison to observational data. We then compare the global properties of the sample to observational data in Section 3 and examine the hot gas profiles of the sample in Section 4. Finally, we discuss our results in Section 5 and present a summary of the main findings in Section 6.

2 NUMERICAL METHOD

This section provides an overview of the cluster sample selection, the model used to resimulate them and how the observable properties were calculated in a manner consistent with observational approaches.

2.1 Sample selection

Due to their limited size, the original EAGLE volumes contain very few clusters, with the largest volume (Ref-L100N1504) containing seven objects with a mass $M_{200} > 10^{14} M_{\odot}$. The C-EAGLE project provides an extension to the cluster environment by performing zoom simulations, which require a population of clusters that a representative sample can be selected from. We use the parent simulation from Barnes et al. (2017) as the basis of our sample selection. It uses a Planck 2013 cosmology (Planck Collaboration et al. 2014) and is

¹ We define M_{200} as the mass enclosed within a sphere of radius r_{200} whose mean density is 200 times the critical density of the Universe.

a cubic periodic volume with a side length of 3.2 Gpc, which is large enough to contain the rarest and most massive haloes expected to form in a Λ CDM cosmology. At $z = 0$, it contains 185 150 haloes with $M_{200} > 10^{14} M_{\odot}$ and 1701 haloes with $M_{200} > 10^{15} M_{\odot}$. The sample was selected by first binning all haloes into 10 evenly spaced log mass bins in the range of $14.0 \leq \log_{10}(M_{200}/M_{\odot}) \leq 15.4$. We did this to ensure that we evenly sampled the chosen mass range, otherwise we would have been biased towards lower masses by the steep slope of the mass function. To ensure that our selected objects would be at the centre of the peak in the local density structure and the focus of our computational resources, we removed objects from the selection bins who had a more massive neighbour within a sphere whose radius was the larger value of either 30 Mpc or $20 r_{200}$. We then randomly picked three haloes from each mass bin to yield a sample of 30 objects, which are listed in Appendix A.

We used the zoom simulation technique (Katz & White 1993; Tormen, Bouchet & White 1997) to resimulate our chosen sample at higher resolution. The Lagrangian region of every cluster was selected so that its volume was devoid of lower resolution particles beyond a cluster-centric radius of at least $5 r_{200}$ at $z = 0$. Additionally, the Lagrangian regions of the *Hydrangea* sample were defined such that they were devoid of lower resolution particles beyond a cluster-centric radius of $10 r_{200}$, enabling studies of galaxy evolution as the environment transitions from isolated field to dense cluster. At $z = 127$, the initial glass-like particle configuration of the high-resolution regions was deformed according to the second-order Lagrangian perturbation theory using the method of Jenkins (2010) and PANPHASIA (Jenkins 2013), a multiscale Gaussian white noise field that is publicly available.² We assumed a flat Λ cold dark matter (Λ CDM) cosmology based on the *Planck* 2013 results combined with baryonic acoustic oscillations, *WMAP* polarization and high-multipole moments experiments (Planck Collaboration et al. 2014). The cosmological parameters were $\Omega_b = 0.04825$, $\Omega_m = 0.307$, $\Omega_{\Lambda} = 0.693$, $h \equiv H_0/(100 \text{ km s}^{-1} \text{ Mpc}^{-1}) = 0.6777$, $\sigma_8 = 0.8288$, $n_s = 0.9611$ and $Y = 0.248$. The resolution of the Lagrangian regions was increased to match the resolution of the *EAGLE* 100 Mpc simulation (Ref-L100N1504). The dark matter particles each had a mass of $m_{\text{DM}} = 9.7 \times 10^6 M_{\odot}$ and the gas particles each had an initial mass of $m_{\text{gas}} = 1.8 \times 10^6 M_{\odot}$ (note no h^{-1}). The proper gravitational softening length for the high-resolution region was set to 2.66 comoving kpc for $z > 2.8$, and then kept fixed at 0.70 physical kpc for $z < 2.8$. The minimum smoothing length of the SPH kernel was set to a 10th of the gravitational softening scale.

2.2 The eagle model

We use the *EAGLE* model to resimulate our selected sample. The *EAGLE* subgrid model is based on the model developed for the OWLS (Schaye et al. 2010) project and also used for the GIMIC (Crain et al. 2009) and COSMO-OWLS (Le Brun et al. 2014) models. The subgrid model, the calibration of its free parameters and its numerical convergence are described in detail in S15 and Crain et al. (2015). The code is a heavily modified version of the *N*-body Tree-PM SPH code P-GADGET-3, which was last described in Springel (2005). The hydrodynamics algorithms are collectively known as ‘ANARCHY’ [see Dalla Vecchia in preparation], appendix A of S15 and Schaller et al. (2015)] and consists of an implementation of the pressure-entropy SPH formalism derived by Hopkins (2013), an artificial viscosity

switch (Cullen & Dehnen 2010), an artificial conductivity switch similar to that of Price (2008), the C^2 smoothing kernel with 58 neighbours (Wendland 1995) and the time-step limiter of Durier & Dalla Vecchia (2012). The subgrid model includes radiative cooling, star formation, stellar evolution, feedback due to stellar winds and supernovae, and the seeding, growth and feedback from BHs. We now briefly describe the subgrid model in more detail.

Net cooling rates are calculated on an element-by-element basis following Wiersma, Schaye & Smith (2009a), under the assumption of an optically thin gas in ionization equilibrium, the presence of the cosmic microwave background and an evolving ultraviolet/X-ray background (Haardt & Madau 2001) from galaxies and quasars. This is done by interpolation tables, computed using *CLOUDY* version 07.02 (Ferland et al. 1998), that are a function of density, temperature and redshift for the 11 elements that were found to be important. During reionization, 2 eV per proton mass is injected to account for enhanced photoheating rates. For hydrogen this occurs instantaneously at $z = 11.5$ and for helium this additional heating is Gaussian distributed in redshift, centred on $z = 3.5$ with a width $\sigma(z) = 0.5$. The latter ensures that the observed thermal history of the intergalactic gas is broadly reproduced (Schaye et al. 2000; Wiersma et al. 2009b).

Star formation is modelled stochastically in a way that, by construction, reproduces the observed Kennicutt–Schmidt relation, as cosmological simulations lack the resolution and physics to properly model the cold interstellar gas phase. It is implemented as a pressure-law (Schaye & Dalla Vecchia 2008), subject to a metallicity-dependent density threshold (Schaye 2004). Gas particles whose density exceeds $n_{\text{H}}(Z) = 10^{-1} \text{ cm}^{-3} (Z/0.002)^{-0.64}$, where Z is the gas metallicity, are eligible to form stars. Lacking the resolution and physics to model the cold gas phase, a temperature floor, $T_{\text{eos}}(\rho_{\text{g}})$, is imposed that corresponds to the equation of state $P_{\text{eos}} \propto \rho_{\text{g}}^{4/3}$, normalized to $T_{\text{eos}} = 8 \times 10^3 \text{ K}$ at $n_{\text{H}} = 10^{-1} \text{ cm}^{-3}$. This helps to prevent spurious fragmentation. Stellar evolution and the resulting chemical enrichment is based upon Wiersma et al. (2009b). Star particles are treated as simple stellar populations with a mass range of $0.1\text{--}100 M_{\odot}$ and a Chabrier (2003) initial mass function. Mass loss and the release of 11 chemical elements due to winds from massive stars, asymptotic giant branch stars and type Ia and type II supernovae are tracked. Feedback from star formation is implemented using the stochastic thermal model of Dalla Vecchia & Schaye (2012). The energy injected heats a particle by a fixed temperature increment, $\Delta T = 10^{7.5} \text{ K}$, to prevent spurious numerical losses. The energy per unit of stellar mass formed, which sets the probability of heating events, depends on the local gas density and metallicity and is calibrated to ensure that the galaxy stellar mass function and galaxy size–mass relation are a good match to the observed relations at $z = 0.1$ (Crain et al. 2015).

Feedback from supermassive BHs is a critical component of structure formation simulations, shaping the bright end of the galaxy luminosity function (e.g. Bower et al. 2006; Croton et al. 2006), the gas content of clusters (e.g. Puchwein et al. 2008; Fabjan et al. 2010; McCarthy et al. 2010) and preventing the onset of the overcooling problem (McCarthy et al. 2011). The prescription for the seeding, growth and feedback from BHs is based on Springel, Di Matteo & Hernquist (2005) with modifications from Booth & Schaye (2009) and Rosas-Guevara et al. (2015). Seed BHs are placed in the centre of every halo with a total mass greater than $10^{10} M_{\odot} h^{-1}$, with the highest density gas particle being converted to a BH particle with a subgrid seed mass of $m_{\text{BH}} = 10^5 M_{\odot} h^{-1}$. Haloes are identified by an on-the-fly friends-of-friends algorithm with a linking length of $b = 0.2$ (Davis et al. 1985). BHs can grow either by merging

² The phase descriptor of the parent volume is given in Appendix A and in table B1 of S15.

with other BHs or via the accretion of gas. The gas accretion rate is Eddington limited and depends on the mass of the BH, the local gas density and temperature, and the relative velocity and angular momentum of the gas compared to the BH:

$$\dot{m}_{\text{accr}} = \dot{m}_{\text{Bondi}} \times \min(C_{\text{visc}}^{-1}(c_s/V_\phi)^3, 1), \quad (1)$$

where \dot{m}_{Bondi} is the Bondi & Hoyle (1944) spherically symmetric accretion rate, c_s is the sound speed of the gas and V_ϕ is the rotation speed of the gas around the BH (see equation 16 of Rosas-Guevara et al. 2015). C_{visc} is a free parameter for the effective viscosity of the subgrid accretion disc/torus and larger values correspond to a lower kinetic viscosity, which delays the growth of BHs and the onset of feedback events. We note that the results are remarkably insensitive to a non-zero value of C_{visc} (Bower et al. 2017). The growth of the BH is then given by $\dot{m}_{\text{BH}} = (1 - \epsilon_r)\dot{m}_{\text{accr}}$, where $\epsilon_r = 0.1$ is the radiative efficiency. The accretion rate is not multiplied by a constant or density-dependent factor as is common to many previous studies (Springel et al. 2005; Booth & Schaye 2009), as at the resolution of EAGLE, we resolve sufficiently high-gas densities and the accretion rate is sufficient for the BH seeds to grow by Bondi–Hoyle accretion.

AGN feedback is implemented in a similar way to stellar feedback, whereby thermal energy is injected stochastically. Energy is injected at a rate of $\epsilon_f \epsilon_r \dot{m}_{\text{accr}} c^2$, where c is the speed of light and $\epsilon_f = 0.15$ is the fraction of energy that couples to the gas. Booth & Schaye (2009) found that for OWLS, this value yielded agreement with the observed BH masses and S15 found that the same holds for EAGLE. In order to prevent spurious numerical losses and enable the feedback to do work on the gas before the injected energy is radiated, the BH stores accretion energy until it reaches a critical energy, at which point it has enough energy to heat n_{heat} particles by a temperature ΔT . In the EAGLE model, $n_{\text{heat}} = 1$. In the pressure–entropy SPH formalism used in EAGLE, the weighted density and entropic function are coupled (S15, see appendix A1.1). For large changes in internal energy, such as AGN feedback events, the iterative scheme used to adjust the density and entropic function fails to adequately conserve energy if many particles are heated simultaneously. This leads to a violation of energy conservation. The probability of heating a neighbour particle is given by

$$P = \frac{E_{\text{BH}}}{\Delta \epsilon_{\text{AGN}} N_{\text{ngb}} \langle m_{\text{gas}} \rangle}, \quad (2)$$

where E_{BH} is the energy reservoir of the BH, $\Delta \epsilon_{\text{AGN}}$ is the specific energy associated with heating a particle by ΔT , N_{ngb} is the number of gas neighbours and $\langle m_{\text{gas}} \rangle$ is their average mass. As E_{BH} increases the probability of heating many neighbours increases and a violation of energy conservation becomes more likely. To prevent this, the probability is capped at a value of $P_{\text{AGN}} = 0.3$, a value that was chosen following testing of the iterative scheme. If the unused energy remains above the critical threshold for a feedback event then the time step of the BH is shortened and the energy spread out over successive steps.

S15 present three calibrated models (REF, AGNdT9 and Recal) that produce a similarly good match to the observed galaxy stellar mass function and galaxy mass–size relation. As we are running at standard EAGLE resolution, we ignore the Recal model, which is relevant for simulations with $8 \times$ higher mass resolution. From figs 15 and 16 of S15, it is clear that the AGNdT9 model provides a better match to the observed gas fraction–total mass and X-ray luminosity–temperature relations of low-mass groups ($M_{500} < 10^{13.5} M_\odot$). Therefore, we select the AGNdT9 model as our fiducial model for the C-EAGLE project; the free parameters of

Table 1. Values of the AGN feedback free parameters of the AGNdT9 model, used for the C-EAGLE simulations and the 50 Mpc EAGLE volume, and the REF model, used for the 100 Mpc EAGLE volume.

Model	n_{heat}	C_{visc}	ΔT (K)
AGNdT9	1	$2\pi \times 10^2$	10^9
REF	1	2π	$10^{8.5}$

the AGN feedback for this model and for the REF model are given in Table 1. This model was calibrated in a 50 Mpc cubic volume to produce good agreement between simulated and observed galaxies. Furthermore, Crain et al. (2015) showed that the simulated BH–stellar mass relation and galaxy mass–size relation are sensitive to the feedback parameters used. We do not recalibrate the model to cluster scale objects and instead choose to retain the match to the observed galaxy properties. Therefore, the properties of the ICM for the C-EAGLE clusters are a prediction of a model that produces reasonably realistic field galaxies.

2.3 Calculating X-ray and Sunyaev–Zel’dovich properties

Inhomogeneities in the hot gas, e.g. the presence of multitemperature structures, can significantly bias the ICM properties inferred from X-ray observations (e.g. Nagai, Vikhlinin & Kravtsov 2007a; Khedekar et al. 2013). Therefore, when comparing simulations to observations it is vital to make a like-with-like comparison and compute properties from simulated clusters in a manner more consistent with observational techniques.

Following Le Brun et al. (2014), we produce mock X-ray spectra for each cluster by first computing a rest-frame X-ray spectrum in the 0.05–100.0 keV band for each gas particle, using their individual density, temperature and SPH-smoothed metallicity. We use the Astrophysical Plasma Emission Code (APEC; Smith et al. 2001) via the PYATOMDB module with atomic data from ATOMDB v3.0.3 (last described in Foster et al. 2012). For each of the 11 elements, we calculate an individual spectrum, ignoring particles with a temperature less than 10^5 K, as they will have negligible X-ray emission.

Particles are then binned, in three-dimensional (3D), into 25 logarithmically spaced radial bins centred on the potential minimum and the spectra are summed for each bin. The binned spectra are scaled by the relative abundance of the heavy elements, as the fiducial spectra assume solar abundances specified by Anders & Grevesse (1989). The energy resolution of a spectrum is 150 eV between 0.05 and 10.0 keV and we use a further 10 logarithmically spaced bins between 10.0 and 100.0 keV. A single temperature, fixed metallicity APEC model is then fitted to the spectrum in the range of 0.5 and 10.0 keV, for each radial bin, to derive an estimate of the density, temperature and metallicity. During the fit we multiply the spectrum by the effective area of *Chandra* for each energy bin to provide a closer match to typical X-ray observations.

We then perform a hydrostatic analysis of each cluster using the X-ray-derived density and temperature profiles. We fit the density and temperature models proposed by Vikhlinin et al. (2006) to obtain a hydrostatic mass profile. We then estimate various cluster masses and radii, such as M_{500} and r_{500} , from the hydrostatic analysis. We calculate properties, such as gas mass, $M_{\text{gas},500}$, or Sunyaev–Zel’dovich (SZ) flux, $Y_{\text{SZ},500}$, by summing the

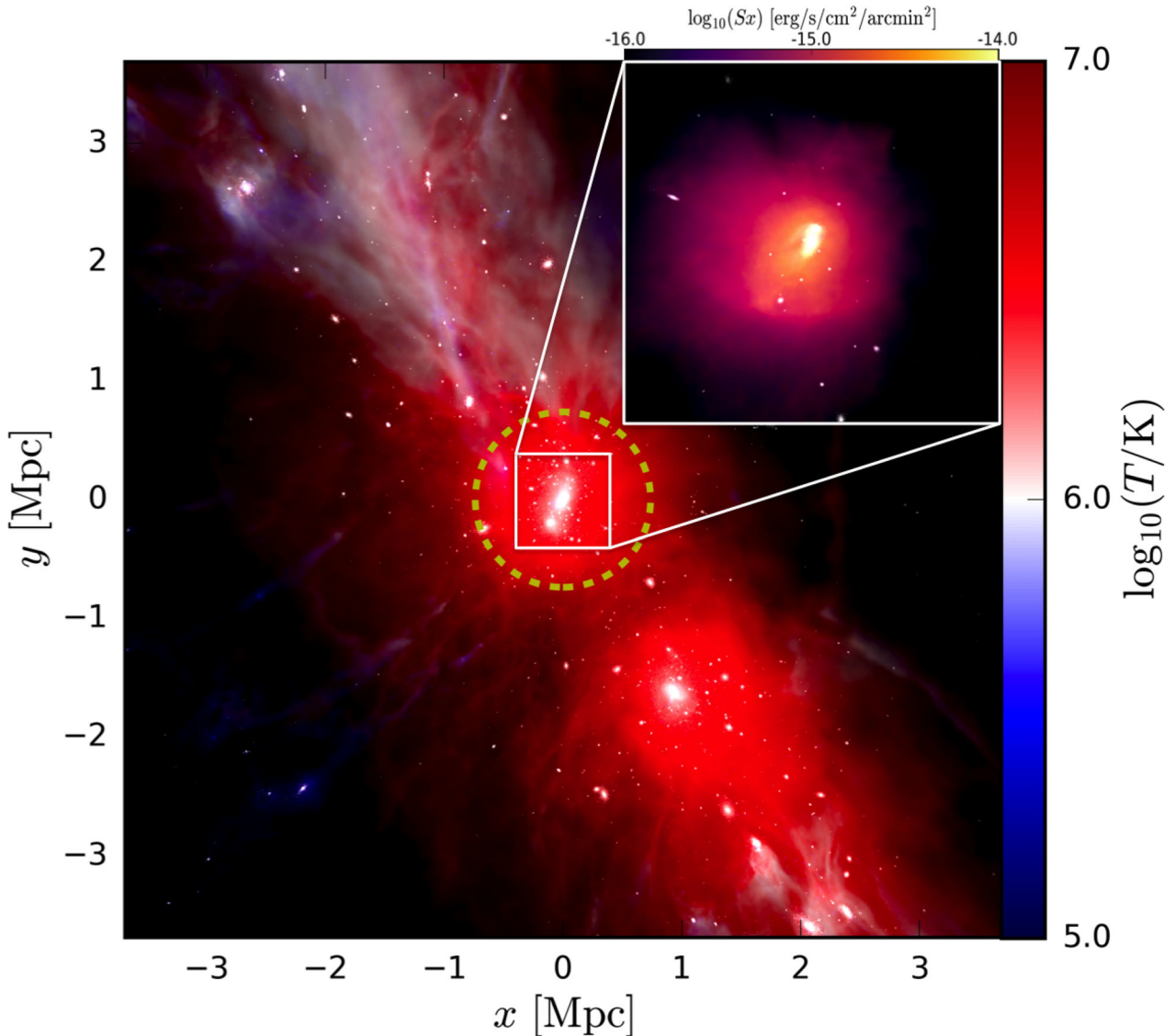


Figure 1. Image of CE-5 and its environment at $z = 0.1$, resimulated using the *EAGLE* AGNdT9 model. The colour map shows the gas, with the intensity depicting the density of the gas and the colour depicting the temperature of the gas. Stellar particles are shown by the white points and the dashed yellow circle denotes r_{200} . The inset colour map shows the X-ray surface brightness from a cubic region of $2 r_{500}$, centred on the cluster’s centre of potential.

properties of the particles that fall within the estimated X-ray apertures. A core-excised quantity is calculated by summing particles that fall in the radial range of 0.15–1.0 of the specified aperture. To calculate X-ray luminosities, we integrate the spectra of particles that fall within the aperture in the required energy band; for example, the soft band luminosities are calculated between 0.5 and 2.0 keV. To estimate a spectroscopic X-ray temperature within an aperture, we sum the spectra of all particles that fall within it and then fit a single temperature *APEC* model to the combined spectrum. All quantities that are derived in this way are labelled as ‘spec’ quantities. In Appendix A, we also provide the values of estimated quantities for each cluster at $z = 0.1$.

In Fig. 1, we show an image of the gas and stars for CE-5 at $z = 0.1$. The resolution and subgrid physics of the *EAGLE* model enable the simulation to capture the formation of galaxies in the dense cluster environment and the surrounding filamentary structures. The inset panel shows soft band (0.5–2.0 keV) X-ray surface brightness within $2 r_{500}$ of the potential minimum of the cluster.

3 GLOBAL PROPERTIES

In this section, we compare the global properties of the *C-EAGLE* sample at $z = 0.1$ with low-redshift ($z \leq 0.25$) observations. We also plot the groups and clusters from the periodic volumes of the *EAGLE* simulations. We label groups and clusters from the 100 Mpc volume run with the *EAGLE* reference model as ‘REF’ and those from the 50 Mpc volume run with the AGNdT9 model as ‘AGNdT9’. To ensure a fair comparison with observational data, we use quantities estimated via the mock X-ray analysis pipeline. However, we stress that estimated masses assume that the cluster is in hydrostatic equilibrium and that the X-ray-estimated density and temperature profiles are good approximations of the true profiles. First, we test this assumption.

3.1 Bias and scatter of estimated masses

To examine the scatter and bias introduced by using estimated masses rather than the true masses, we plot the ratio of the

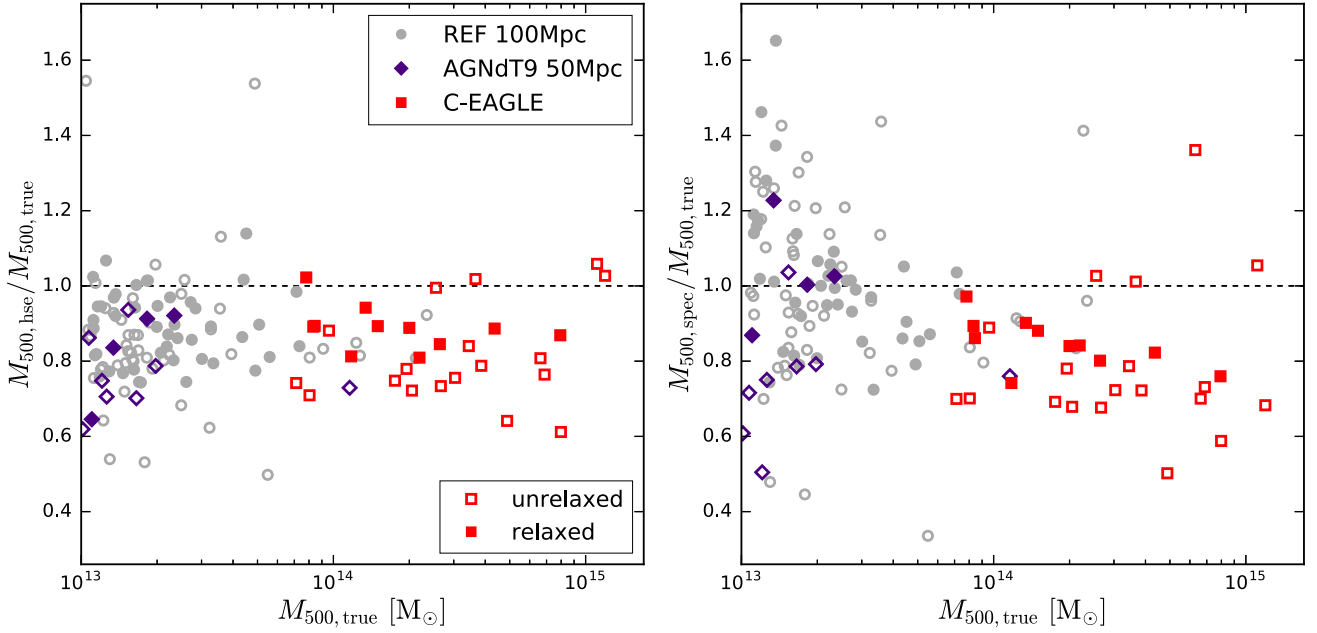


Figure 2. Ratio of estimated to true mass as a function of true mass at $z = 0.1$ for the C-EAGLE clusters (red squares), as well as the groups and clusters from the EAGLE REF (grey circles) and AGNdT9 models (purple diamonds). The left-hand panel shows hydrostatic mass estimates, calculated by fitting Vikhlinin et al. (2006) models to the true profiles, while the right-hand panel shows X-ray spectroscopic mass estimates, calculated by fitting Vikhlinin et al. (2006) models to profiles estimated from the mock X-ray pipeline. Clusters that are defined as relaxed (unrelaxed) are shown by the filled (open) points. The dashed black line indicates no bias.

estimated M_{500} over $M_{500,true}$ as a function of $M_{500,true}$, where the ‘true’ M_{500} values are calculated via summation of particle masses that fall within the true r_{500} . To separate the effects of assuming hydrostatic equilibrium and estimating the profiles in an observational manner via the X-ray pipeline, we also fit the Vikhlinin et al. (2006) models to the true density and temperature profiles and label any quantity computed in this manner as ‘hse’.

Fig. 2 shows the scatter and bias introduced by estimating the mass for the C-EAGLE clusters as well as the REF and AGNdT9 groups and clusters. The assumption of hydrostatic equilibrium introduces a median bias $b_{hse} = 0.16 \pm 0.04$ for the C-EAGLE clusters, where $b_{hse} = 1 - M_{500,hse}/M_{500,true}$ and the error is calculated by bootstrap resampling the data 10 000 times. The assumption of hydrostatic equilibrium introduces a similar bias in the REF and AGNdT9 groups and clusters, which yield values of $b_{hse} = 0.14 \pm 0.02$ and $b_{hse} = 0.21 \pm 0.06$, respectively. We find that the bias introduced by assuming hydrostatic equilibrium is independent of mass. This is consistent with previous simulation work that has analysed true profiles (Nelson et al. 2014; Biffi et al. 2016; Henson et al. 2017). For the C-EAGLE clusters, the profiles derived from the mock X-ray analysis produce a slightly larger bias with a median value $b_{spec} = 0.22 \pm 0.04$. Combined with the REF and AGNdT9 groups and clusters, it is clear that using the X-ray derived profiles increases the scatter in the estimated masses, with C-EAGLE yielding rms values of $\sigma_{hse} = 0.16$ and $\sigma_{spec} = 0.18$ for the hse and spec values, respectively. Additionally, the spectroscopic bias appears to become mildly mass dependent for the mock X-ray pipeline mass estimates. This is consistent with previous simulation work that has made use of mock X-ray pipelines; Le Brun et al. (2014) saw an increase in scatter for low-mass ($< 10^{13} M_{\odot}$) groups and a median bias consistent with zero and Henson et al. (2017) found that the bias from mock X-ray pipelines showed a mild mass dependence.

Several of the C-EAGLE clusters have mass estimates that are more than 30 per cent discrepant from their true mass. To better

understand the impact of the dynamical state of a cluster on the estimation of its mass, we classify each object as relaxed or unrelaxed. Theoretically, there are many ways of defining whether a cluster is relaxed (see Neto et al. 2007; Duffy et al. 2008; Klypin, Trujillo-Gomez & Primack 2011; Dutton & Macciò 2014; Klypin et al. 2016). In this work, we define a cluster as being relaxed if

$$E_{kin,500,spec}/E_{therm,500,spec} < 0.1,$$

where $E_{kin,500,spec}$ is the sum of the kinetic energy of the gas particles, with the bulk motion of the cluster removed, inside r_{500} and $E_{therm,500,spec}$ is the sum of the thermal energy of the gas particles within r_{500} . In all figures, we denote relaxed (unrelaxed) clusters by solid (open) points. Using this criterion, 11 of the 30 C-EAGLE clusters are defined as relaxed.

Selecting only relaxed (unrelaxed) C-EAGLE clusters, we measure median mass biases $b_{hse} = 0.14 \pm 0.02$ (0.19 ± 0.03) and $b_{spec} = 0.16 \pm 0.03$ (0.26 ± 0.04). Thus, the mass estimates of relaxed clusters show a small decrease in the level of bias compared to the full sample, but the scatter in the mass estimate reduces significantly to $\sigma_{hse} = 0.06$ and $\sigma_{spec} = 0.06$. The bias of unrelaxed clusters increases slightly and the scatter about the median value increases to $\sigma_{hse} = 0.23$ and $\sigma_{spec} = 0.21$. We see a similar trend for the REF and AGNdT9 groups and clusters. We also find that all C-EAGLE clusters with a mass estimate that is more than 30 per cent discrepant from its true mass have $E_{kin,500,spec}/E_{therm,500,spec} > 0.1$, demonstrating the impact of the dynamical state of the cluster on its estimated mass.

3.2 Gas, stellar and BH masses

Theoretical work has shown that the gas and stellar content of clusters is largely controlled by stellar and particularly AGN feedback (Voit et al. 2003; Bower et al. 2008; Fabjan et al. 2010; McCarthy et al. 2011; Planelles et al. 2013; Le Brun et al. 2014; Pike et al. 2014;

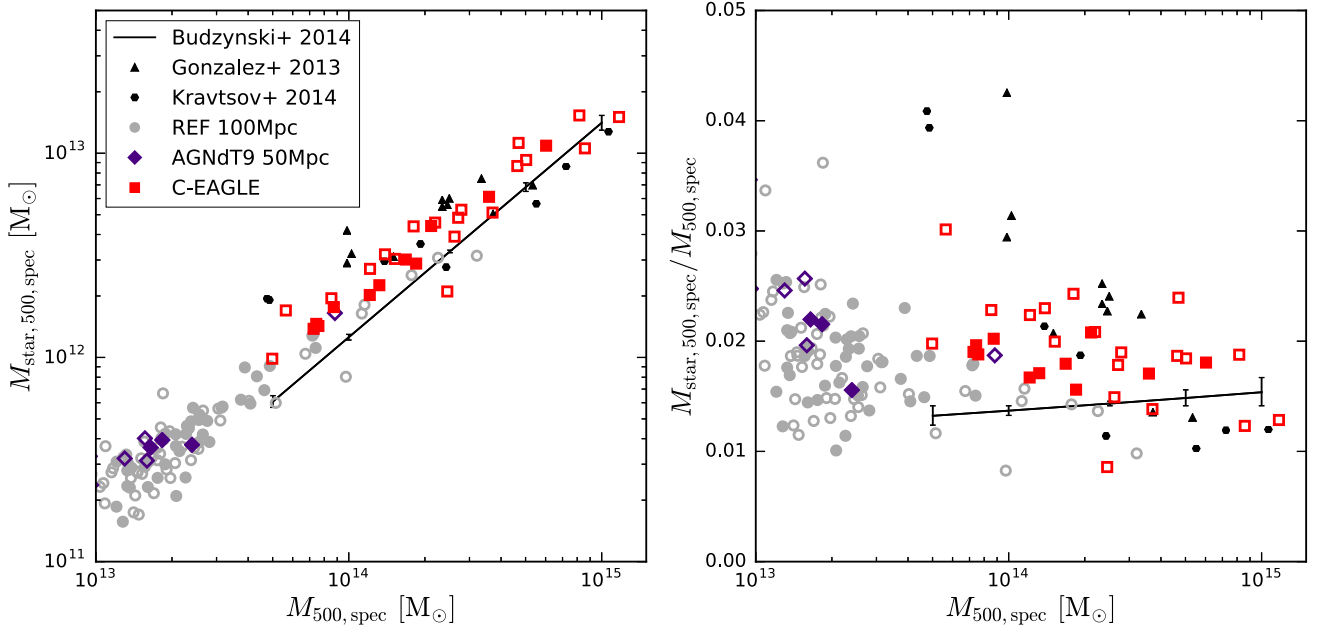


Figure 3. Integrated stellar mass (left-hand panel) and stellar fraction (right-hand panel) within $r_{500,\text{spec}}$ as a function of estimated total mass at $z = 0.1$ for the C-EAGLE clusters and the REF and AGNdT9 groups and clusters. Marker styles are the same as in Fig. 2. The black triangles and hexagons show the observational data from Gonzalez et al. (2013) and Kravtsov et al. (2014), respectively, and the black line with error bars shows the best-fitting result from Budzynski et al. (2014) from stacking SDSS images.

Hahn et al. 2017; McCarthy et al. 2017). Therefore, the gas, stellar and BH masses of the C-EAGLE clusters provide a test, orthogonal to usual tests of galaxy formation, of the feedback model and whether calibration on against galaxy properties alone leads to a reasonably realistic ICM.

3.2.1 Stellar mass

We plot the integrated stellar mass and stellar fraction within $r_{500,\text{spec}}$ as a function of the estimated total mass at $z = 0.1$ in the left-hand and right-hand panels of Fig. 3, respectively. To ensure a fair comparison, we only compare the C-EAGLE, REF and AGNdT9 samples against observations where the mass of the system has been estimated via high-quality X-ray observations. The stellar mass–total mass relation within $r_{500,\text{true}}$ is shown in fig. 4 of Bahé et al. (2017). All of the observations and the simulations include the intracluster light in the stellar mass estimate and assume a Chabrier (2003) initial mass function. First, we note that the observations do not appear to be consistent with each other. The results of Budzynski et al. (2014) have a lower normalization compared to the results of Kravtsov, Vikhlinin & Meshcheryakov (2014) and Gonzalez et al. (2013) at $M_{500} = 10^{14} M_{\odot}$, and the stellar fraction of Budzynski et al. (2014) increases slightly with total mass while Kravtsov et al. (2014) and Gonzalez et al. (2013) show a strongly decreasing stellar fraction with increasing total mass. This is most likely due to the different selection criteria and methods used in the observations.

The C-EAGLE clusters provide a consistent extension into the high-mass regime of the original periodic volumes. Estimating the cluster’s mass from the X-ray hydrostatic analysis leads to an increased scatter about the relation compared to using the true mass, increasing from $\sigma_{\log_{10},\text{true}} = 0.07$ to $\sigma_{\log_{10},\text{spec}} = 0.10$.³ If only relaxed systems

are considered the scatter reduces substantially to $\sigma_{\log_{10},\text{spec}} = 0.03$. Taken together with groups and clusters from the periodic volumes, the C-EAGLE clusters show reasonable agreement with the observations. They reproduce the trend of the best-fitting stellar mass–total mass relation of Budzynski et al. (2014), but have a slightly higher normalization. In contrast, they reproduce the normalization of the Gonzalez et al. (2013) and Kravtsov et al. (2014) observations and are within the intrinsic scatter, but the trend with total mass is different. This is better seen in the stellar fraction–total mass relation, where the results of Gonzalez et al. (2013) and Kravtsov et al. (2014) show an increase in stellar fraction for lower mass objects, while the C-EAGLE clusters combined with the REF and AGNdT9 objects produce a roughly constant stellar mass fraction over two decades in mass. Overall, ≈ 2 per cent of the total mass of the simulated groups and clusters consists of stars. This is consistent with previous numerical work, which has shown that the stars make up a few per cent of the cluster mass and the trend with mass is relatively mild (e.g. Planelles et al. 2013; Pike et al. 2014; Hahn et al. 2017). The EAGLE model was calibrated to reproduce the observed galaxy stellar mass function of the field at low redshift, but dense cluster environments were not present in the calibration volumes due to their limited size. The level of agreement between the C-EAGLE clusters and the observations is, therefore, reassuring and demonstrates that the galaxy calibrations continue to work in the cluster regime.

3.2.2 BH masses

We examine the properties of the supermassive BHs that form in the C-EAGLE clusters in Fig. 4. In the left-hand panel, we plot BH mass as a function of stellar mass, calculated within a 3D aperture of radius 50 kpc, for those BHs with a mass $> 10^7 M_{\odot}$ that fall within $r_{200,\text{true}}$.

³ We measure the scatter by fitting a power law to the stellar mass–total mass relation for the full C-EAGLE sample, the relaxed subset and the unrelaxed

subset. Further details of our method and a summary of the results for all scaling relations can be found in Appendix B.

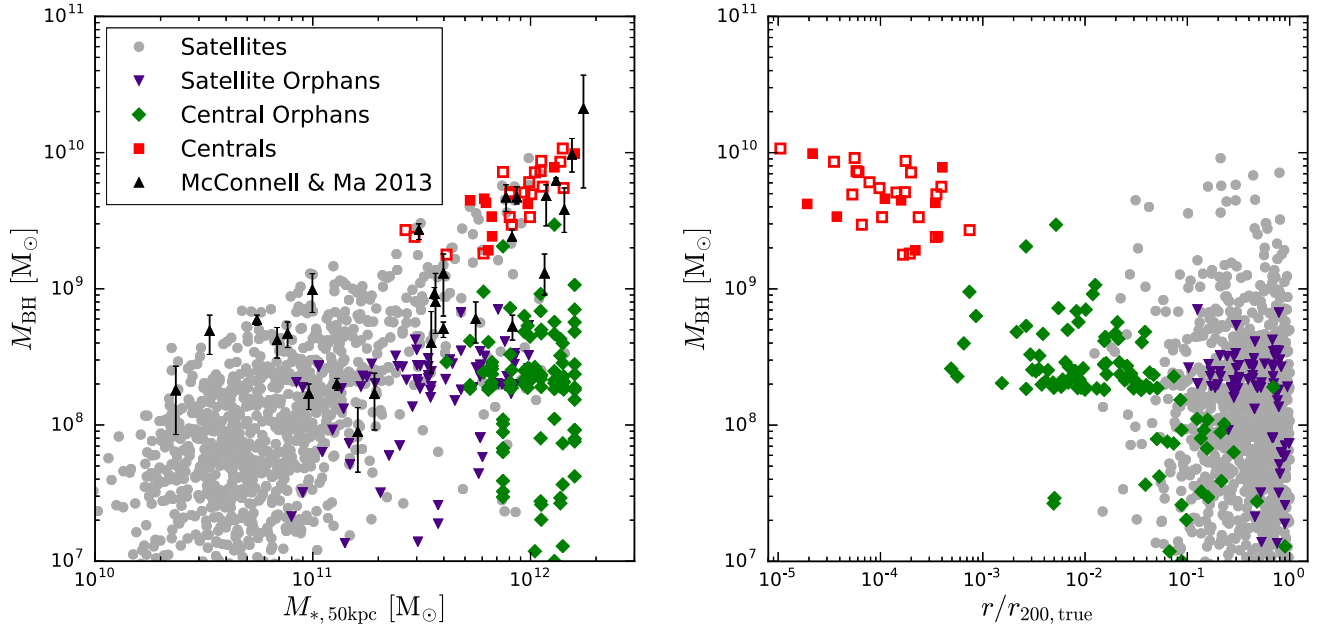


Figure 4. BH mass as a function of stellar mass within a 3D 50 kpc aperture (left-hand panel) and cluster-centric radius (right panel) for those BHs that fall within r_{200} of a C-EAGLE cluster at $z = 0.1$. The population is divided between centrals (red squares), satellites (grey circles), orphan BHs associated with the main halo (green diamonds) and orphans associated with satellites (purple triangles). For the centrals, we denote whether the cluster is relaxed (unrelaxed) by filled (open) marker style. The black triangles show the observations of McConnell & Ma (2013) for early type galaxies.

In total, the C-EAGLE sample contains 1358 BHs within $r_{200,\text{true}}$. The aperture radius is a choice and we select 50 kpc as it mimics common observational choices. We decompose the bound population into centrals (those bound to the main halo) and satellites (those bound to subhaloes) as determined by the SUBFIND algorithm (Springel et al. 2001; Dolag et al. 2009). We find that 90 per cent of all BHs are defined as satellites. The central and satellite BHs form a continuous population on the BH mass–stellar mass relation. We compare the simulation relation to the observed early type galaxies from the compilation of McConnell & Ma (2013). We find good agreement with the observed relation, with the C-EAGLE clusters reproducing both the observed trend and normalization. As demonstrated in Booth & Schaye (2010), the normalization of the BH mass–stellar mass relation is set by the feedback efficiency and the chosen value of $\epsilon_f = 0.15$ has been shown to work in low-resolution simulations of galaxies (Booth & Schaye 2009), clusters (Le Brun et al. 2014) and in the EAGLE periodic volumes (S15).

Additionally, we define orphan BHs as those that are bound to either the main halo or a subhalo, but are not the most massive BH associated with it. We find that ≈ 15 per cent of BHs within $r_{200,\text{true}}$ are classified as orphans. We further classify the orphans as being associated either with the main halo or subhaloes and find around 50 per cent of the orphans are linked to the main halo. The number of orphans increases strongly with cluster mass and only subhaloes with a stellar mass $> 10^{11} M_\odot$ contain orphans. The EAGLE model allows BHs whose mass exceeds $100m_{\text{gas}}$ to become dynamically independent and limits the distance that BHs less massive than this limit can drift towards the potential minimum at each time step. Therefore, when two haloes merge their BHs do not necessarily also instantaneously merge. In the right-hand panel of Fig. 4, we plot BH mass as a function of $r/r_{200,\text{true}}$. The central BHs all lie within 0.1 per cent of $r_{200,\text{true}}$ from the potential minimum, even though they are all dynamically independent. No satellite BH is closer than 1 per cent of $r_{200,\text{true}}$. We find that 43

orphans lie within 1 per cent of $r_{200,\text{true}}$, which equates to a distance of 10–20 kpc from the potential minimum. Therefore, the C-EAGLE simulations predict that there is a population of BHs with a mass $> 10^8 M_\odot$ inside the brightest cluster galaxies (BCGs) that are not the central BH. The orphans, both centrals and satellites, are likely the remnants of previous galaxy mergers, but we note that some may be identified with missed subhaloes as halo finders, such as SUBFIND, have issues identifying subhaloes in dense backgrounds (Muldrew, Pearce & Power 2011). We leave a more detailed examination of the origins and properties of this orphan population to a future study.

3.2.3 Gas mass

In Fig. 5, we plot the gas mass and the gas fraction within $r_{500,\text{spec}}$ as a function of the estimated total mass at $z = 0.1$, in the left-hand and right-hand panels, respectively. Again, we compare against the groups and clusters from the REF and AGNdt9 volumes and observational data. The C-EAGLE clusters provide a consistent extension to the objects from the periodic volumes. The C-EAGLE clusters reproduce the observed trend with halo mass, but they are too gas rich and lie along the top of the observed scatter of the gas mass–total mass relation. Although we selected the AGNdt9 EAGLE model, as it produced a better match to the observed gas fractions of low-mass ($< 10^{13.5} M_\odot$) groups, the subgrid model was not calibrated at cluster scales. Therefore, the gas mass and fractions are a prediction of a model calibrated for field galaxy formation. Many previous numerical simulations yield a better match to the observed gas fraction relation (Planelles et al. 2013; Le Brun et al. 2014; Pike et al. 2014; McCarthy et al. 2017). However, these simulations have significantly lower resolution and use traditional SPH, which has been shown by Schaller et al. (2015) to have an impact on group scales. We note that one of the C-EAGLE clusters is undergoing a

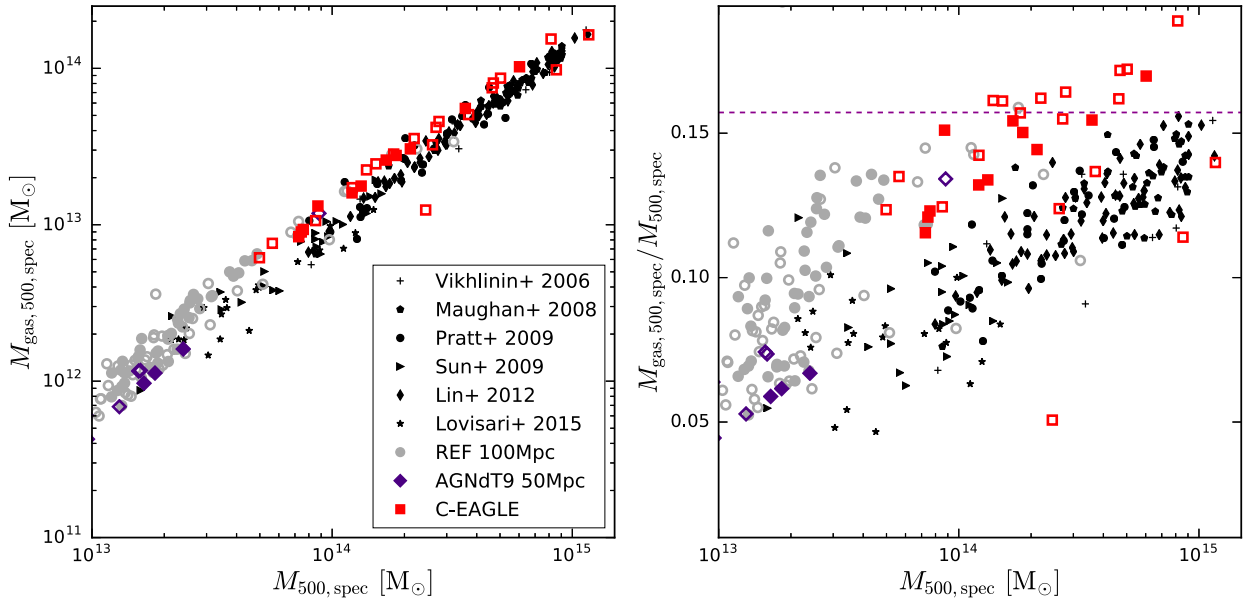


Figure 5. Gas mass as a function of estimated total mass within $r_{500,\text{spec}}$ at $z = 0.1$ for the C-EAGLE clusters and the REF and AGNdT9 groups and clusters. Marker styles are the same as in Fig. 2. The black pluses, pentagons, circles, right-facing triangles, thin diamonds and stars show the observational data from Vikhlinin et al. (2006), Maughan et al. (2008), Pratt et al. (2009), Sun et al. (2009), Lin et al. (2012) and Lovisari et al. (2015), respectively. The left-hand panel shows the gas mass, whereas the right-hand panel shows gas fractions with the magenta dashed line showing the universal baryon fraction, $\Omega_b/\Omega_M = 0.157$.

major merger at $z = 0.1$, displacing the gas from the potential minimum and leading to an estimated gas mass well below the observed relation.

We also plot the gas fractions as the dynamic range of the gas mass–total mass relation hides some of the discrepancies. As shown in S15, the low-mass ($< 10^{13.5} M_\odot$) groups of the periodic volumes show a significant difference in the gas fraction for the two AGN calibrations, with the increased AGN heating temperature of the AGNdT9 model producing gas fractions that are systematically lower by ≈ 30 per cent. However, at cluster scales ($> 10^{14} M_\odot$), the runs with different heating temperatures converge towards to the same relation, although we currently suffer from a small sample size. Gas fractions for some of the C-EAGLE clusters appear to be greater than the universal baryon fraction, $f_b \equiv \Omega_b/\Omega_M = 0.157$. This is due to using an X-ray estimated mass that is biased low compared to the true mass.

To remove the impact of using an estimated mass, we plot the true gas fraction as a function of $M_{500,\text{true}}$ in Fig. 6. The use of true mass significantly reduces the scatter in the relation and all of the simulated clusters are now below the universal fraction, with the most massive clusters reaching ≈ 90 per cent of the universal fraction. However, the simulated clusters and groups above a mass of $M_{500,\text{spec}} \geq 10^{13.5} M_\odot$ are still too gas rich. Previous work has shown that the majority of gas expulsion occurs in the progenitors of a halo at high redshift (McCarthy et al. 2011). These progenitors form at an earlier epoch for more massive objects, when the Universe was denser, making their potentials deeper. The higher-than-observed gas fractions of the groups and clusters suggest that the AGN feedback in the EAGLE model is not efficient enough at removing gas from the deeper potentials of the progenitors of massive groups and clusters ($> 10^{14} M_\odot$). This also leads to some overcooling and larger than observed BCGs (see the right-hand panel of fig. 4 of Bahé et al. 2017). We note that the true gas fractions for rich clusters ($10^{15} M_\odot$) appear to be in agreement with the observational data; however, this is clearly no longer a fair comparison.

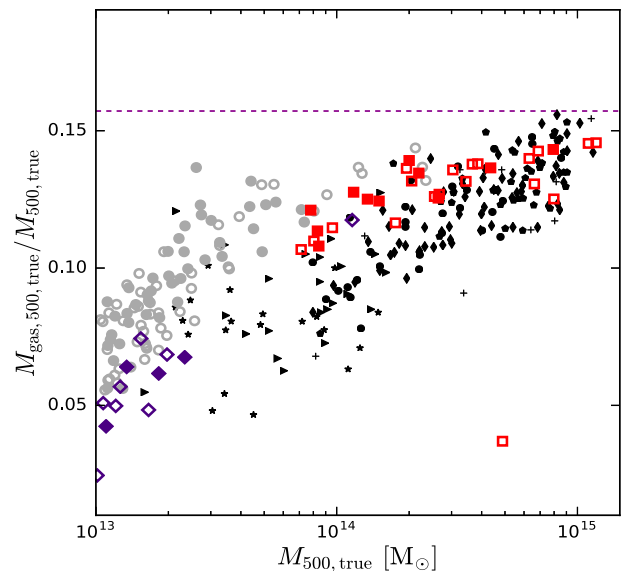


Figure 6. True gas fraction as a function of true total mass, both within $r_{500,\text{spec}}$, at $z = 0.1$ for the C-EAGLE clusters and the REF and AGNdT9 groups and clusters. Marker styles and observations are the same as in Fig. 5.

3.3 X-ray and SZ properties

Having shown that the C-EAGLE clusters have stellar and BH masses in reasonable agreement with the observed relations but are too gas rich, we now examine the X-ray and SZ observable properties at $z = 0.1$ produced by the mock X-ray pipeline.

3.3.1 Spectroscopic temperature–total mass

In the left-hand panel of Fig. 7, we plot the spectroscopic X-ray temperature, $k_B T_{500,\text{spec}}^X$, measured within $r_{500,\text{spec}}$ as a function of

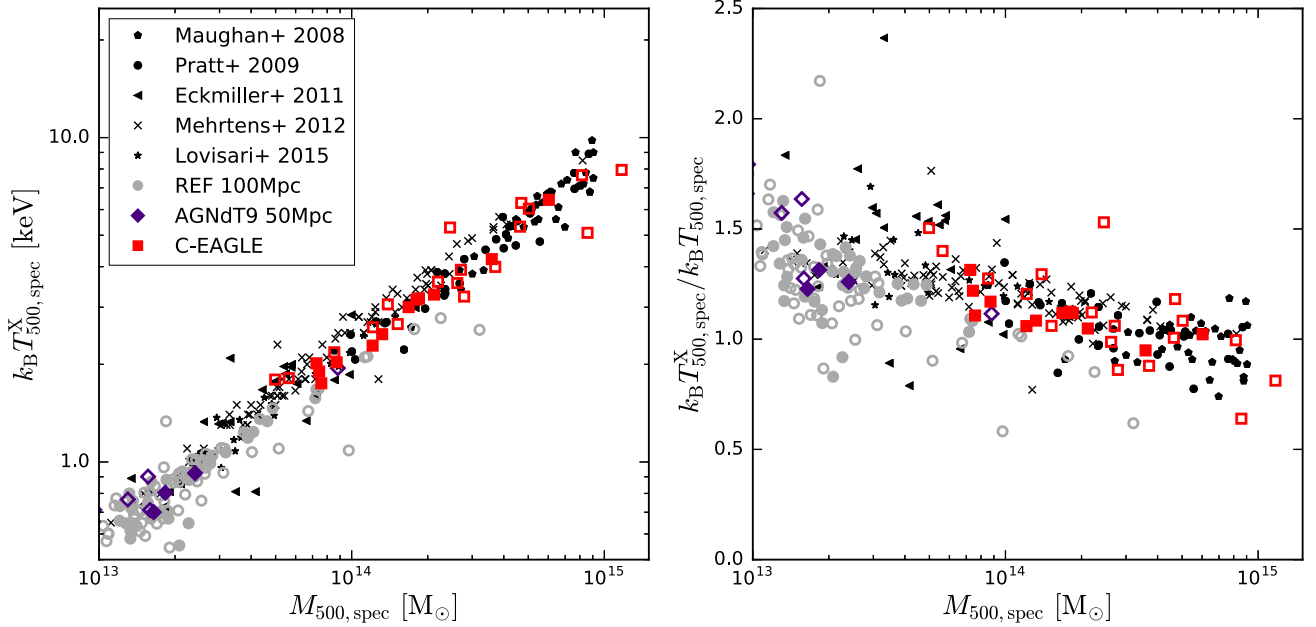


Figure 7. Spectroscopic temperature measured within $r_{500,\text{spec}}$ as a function of estimated total mass at $z = 0.1$ for the *c-EAGLE* clusters and the REF and AGNdt9 groups and clusters. Marker styles are the same as in Fig. 2. The black pentagons, circles, left-facing triangles, crosses and stars show observational data from Maughan et al. (2008), Pratt et al. (2009), Eckmiller et al. (2011), Mehrrens et al. (2012) and Lovisari et al. (2015), respectively. The left-hand panel shows the temperature and the right-hand panel shows the temperature normalized by the virial temperature to approximately remove the mass dependence.

the estimated total mass. The *c-EAGLE* clusters and the groups and clusters from the REF and AGNdt9 volumes are a good match to the observational data, yielding a linear power-law relation over two decades in mass. The *c-EAGLE* simulations have a scatter of $\sigma_{\log_{10}} = 0.06$, which is similar to the value of $\sigma = 0.05$ found by Eckmiller, Hudson & Reiprich (2011) for the HIFLUGCS sample. Though we calculate properties in a manner consistent with observational techniques, we do not make any attempt to account for observational selection effects. The clusters with the largest scatter are all unrelaxed and these objects may not be selected in X-ray samples due to the n_{H}^2 dependence of the emission, which results in a bias towards more relaxed systems.

In the right-hand panel of Fig. 7, we have removed the expected mass dependence of the relation that results from hydrostatic equilibrium by dividing each system by its virial temperature, which we calculate as

$$k_{\text{B}} T_{500,\text{spec}} \equiv G M_{500,\text{spec}} \mu m_{\text{p}} / 2 r_{500,\text{spec}}, \quad (3)$$

where k_{B} is the Boltzmann constant, G is Newton’s gravitational constant, $\mu = 0.59$ is the mean molecular weight of the gas and m_{p} is the proton mass. This allows us to examine the impact of non-gravitational processes, such as cooling and feedback.

The most massive *c-EAGLE* clusters have temperature ratios that are close to one, showing that their temperatures are close to the expected virial values. As the mass of the object decreases the temperature increases relative to the virial temperature, being 50 per cent higher than expected for the $10^{13} M_{\odot}$ groups in the REF and AGNdt9 samples. This demonstrates that the feedback is able to heat and expel gas more efficiently in lower mass objects. This is consistent with previous simulation work (McCarthy et al. 2010; Short et al. 2010; Le Brun et al. 2014; Planelles et al. 2014; Hahn et al. 2017). However, in Le Brun et al. (2014), an increase in the AGN heating temperature produced a flatter normalized temperature relation (see the right hand panel in their Fig. 2), a difference

that we do not see between the REF and AGNdt9 models. This may be due to resolution as the *c-EAGLE* clusters achieve a factor $500 \times$ better mass resolution, which reduces the energy per AGN feedback event and may reduce the impact of feedback on the cluster properties.

3.3.2 X-ray luminosity–total mass

We plot the soft band (0.5–2.0 keV) X-ray luminosity, $L_{\text{X},500,\text{spec}}^{0.5-2.0\text{keV}}$, within $r_{500,\text{spec}}$ as a function of the estimated total mass at $z = 0.1$ in Fig. 8. The *c-EAGLE* clusters show reasonable agreement with the observational data, reproducing the observed trend. The normalization of the simulated trend is marginally high compared to the observed relation, but still within the scatter, and this is due to the clusters being too gas rich. We note that the merging cluster has a significantly decreased X-ray luminosity for its mass. The complete *c-EAGLE* sample has a scatter of $\sigma_{\log_{10}} = 0.30$, which is larger than the values of $\sigma_{\log_{10}} = 0.17$ and $\sigma_{\log_{10}} = 0.25$ found for the REXCESS (Pratt et al. 2009) and HIFLUGCS (Lovisari, Reiprich & Schellenberger 2015) samples, respectively. However, we stress that we do not attempt to account for selection effects. If we select only relaxed clusters the scatter reduces to $\sigma_{\log_{10}} = 0.11$. The *c-EAGLE* clusters are consistent with the clusters of the REF and AGNdt9 samples, with the models producing negligible difference for the X-ray luminosity–mass relation.

At group scales, we find a mild break in the power-law relation. The *EAGLE* model reproduces the observed luminosity–mass relation over two decades in mass and four decades in luminosity. Previous numerical works that include AGN feedback and reproduce the observed gas fraction relation also reproduce the observed X-ray luminosity–total mass relation (e.g. Le Brun et al. 2014; Pike et al. 2014; Planelles et al. 2014; McCarthy et al. 2017), while those without, or with ineffective, AGN feedback (Biffi et al. 2014; Hahn et al. 2017) tend to overestimate the normalization of the relation.

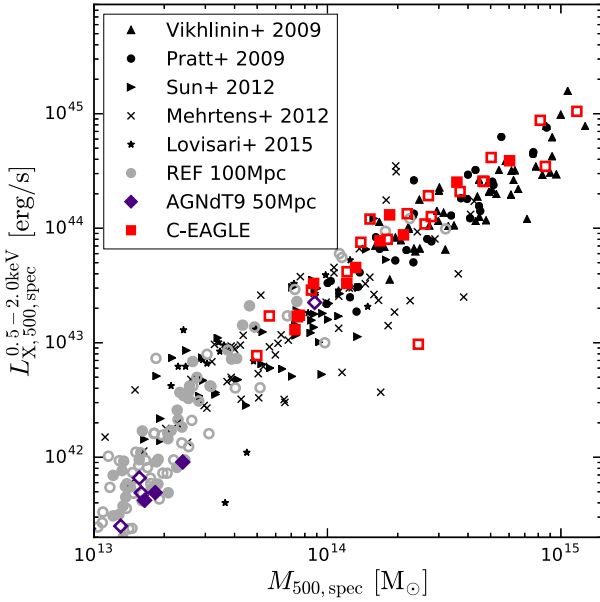


Figure 8. Soft band X-ray luminosity within $r_{500, \text{spec}}$ as a function of estimated total mass at $z = 0.1$ for the C-EAGLE clusters and the REF and AGNdT9 groups and clusters. Marker styles are the same as in Fig. 2. The black triangles, circles, right-facing triangle, crosses and stars show observational data from Vikhlinin et al. (2009), Pratt et al. (2009), Sun (2012), Mehrstens et al. (2012) and Lovisari et al. (2015), respectively.

We note that the scatter increases for observed low-mass groups, but the scatter in the simulated systems does not. However, this increase in the scatter of observed groups may simply reflect the increasing challenge of measuring the total X-ray emission from low-mass groups, rather than a failure of the simulations to reproduce the scatter of the observed relation.

3.3.3 Metallicity–spectroscopic temperature

In Fig. 9, we plot the ‘mass-weighted’ iron abundance within $r_{500, \text{spec}}$ as a function of the spectroscopic temperature at $z = 0.1$. Following Yates, Thomas & Henriques (2017), we estimate the mass-weighted iron abundance via

$$Z_{\text{Fe}, 500, \text{spec}} = \frac{\int_0^{r_{500, \text{spec}}} Z_{\text{Fe}, \text{spec}} \rho_{\text{gas}, \text{spec}} r^2 dr}{\int_0^{r_{500, \text{spec}}} \rho_{\text{gas}, \text{spec}} r^2 dr}, \quad (4)$$

where $Z_{\text{Fe}, \text{spec}}$ and $\rho_{\text{gas}, \text{spec}}$ are the iron abundance and gas density profiles, respectively, estimated from single temperature fits to the summed particle spectra for 25 3D cluster-centric radial bins and we integrate over those bins that fall within $r_{500, \text{spec}}$. We compare the simulated samples against consolidated data taken from Yates et al. (2017); see the references therein for more details on the homogenization process and the different samples. If more than one estimate for the observed mass-weighted metallicity or temperature was present for a system, the mean value was taken. We have scaled both the simulated and observational data to the solar abundances of Asplund et al. (2009).

The C-EAGLE clusters show reasonable agreement with the observed metallicity–temperature relation, with a median metallicity of $Z_{\text{Fe}, 500, \text{spec}}^{\text{med}} = 0.23 Z_{\odot}$ and an rms scatter of $\sigma = 0.09$. The REF and AGNdT9 group clusters are also in agreement with the data, with median metallicities of $Z_{\text{Fe}, 500, \text{spec}}^{\text{med}} = 0.26 Z_{\odot}$ and $Z_{\text{Fe}, 500, \text{spec}}^{\text{med}} = 0.22 Z_{\odot}$ and scatters of $\sigma = 0.07$ and $\sigma = 0.06$, respectively. Overall, the EAGLE model produces a relatively flat

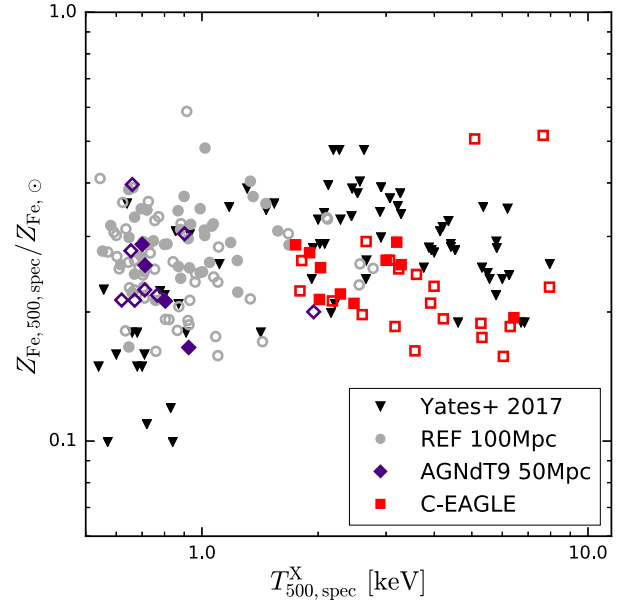


Figure 9. Mass-weighted iron abundance measured within $r_{500, \text{spec}}$ as a function of spectroscopic temperature at $z = 0.1$ for the C-EAGLE clusters and the REF and AGNdT9 groups and clusters. Marker styles are the same as in Fig. 2. The black triangles show the homogenized observational data from Yates et al. (2017).

mass-weighted metallicity–spectroscopic temperature relation from low-mass groups to rich clusters. Reproducing the global metallicities of groups and clusters has been a challenge for previous simulation works, with Martizzi et al. (2016) underpredicting the total metal content of clusters, Planelles et al. (2014) overpredicting the metallicity of more massive clusters and Yates et al. (2017) overpredicting the metallicity of groups. Recent observational (McDonald et al. 2016; Mantz et al. 2017) and numerical (Biffi et al. 2017) results demonstrate that the global cluster metallicity shows very little evolution, suggesting that simulations must carefully model the early galaxy formation processes to reproduce the observed cluster metallicity.

3.3.4 Y-total mass relations

The SZ signal, Y_{SZ} , is a measure of the total thermal energy content of the ICM and is thought to be relatively insensitive to the details of the baryonic physics within the cluster volume (da Silva et al. 2004; Nagai 2006), although see Le Brun et al. (2014). The X-ray analogue of the SZ signal, Y_X , was first proposed by Kravtsov, Vikhlinin & Nagai (2006) and we define it as the product of the core-excised spectroscopic temperature and the gas mass. The difference between the two quantities is that Y_{SZ} is dependent on the mass-weighted temperature, whereas Y_X is dependent on the core-excised spectroscopic temperature. We plot Y_{SZ} within $5r_{500, \text{spec}}$ and Y_X within $r_{500, \text{spec}}$ as a function of the estimated total mass within $r_{500, \text{spec}}$ at $z = 0.1$ in Fig. 10. For the SZ signal, we compare against clusters from the second *Planck* SZ catalogue (Planck Collaboration et al. 2016) that are not infrared contaminated, have a neural quality > 0.4 (the recommended quality threshold), have a mass estimate and are at a redshift of $z < 0.25$. This yields an observed sample of over 600 clusters and we remove the redshift dependence of the fluxes by scaling by the square of the angular diameter distance and applying a self-similar scaling of $E^{-2/3}(z)$, where $E(z) \equiv H(z)/H_0 = \sqrt{\Omega_M(1+z)^3 + \Omega_\Lambda}$. We then binned

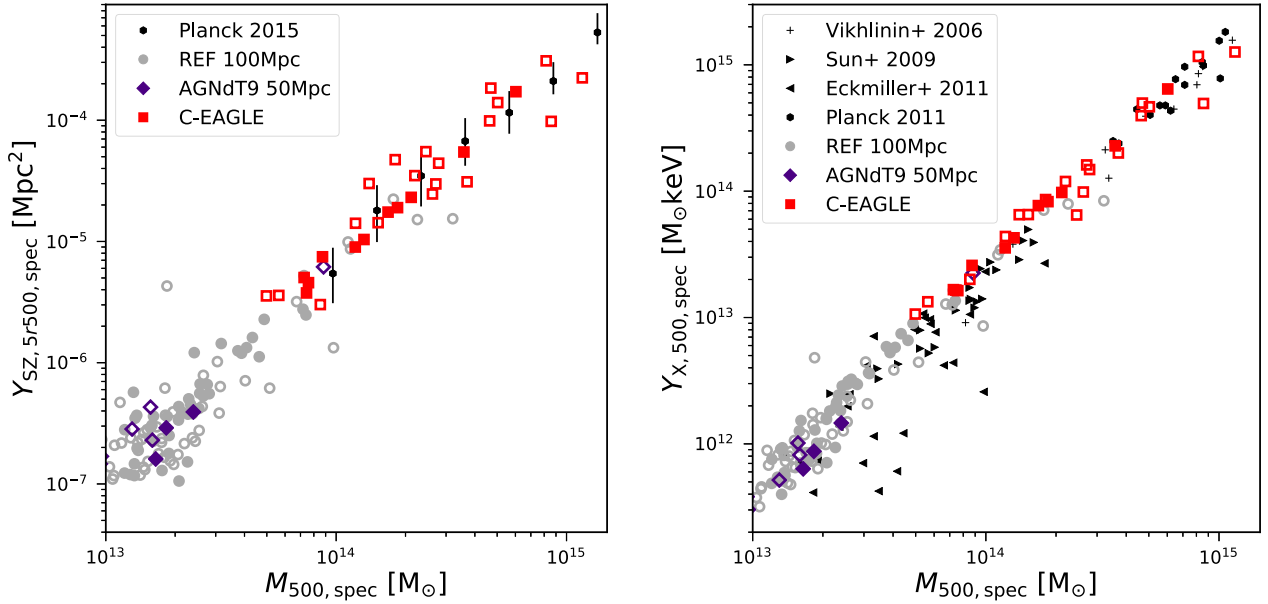


Figure 10. SZ signal within $5r_{500,\text{spec}}$ (left-hand panel) and Y_X within $r_{500,\text{spec}}$ (right-hand panel) as a function of estimated total mass within $r_{500,\text{spec}}$ at $z = 0.1$ for the C-EAGLE clusters and the REF and AGNdt9 groups and clusters. Marker styles are the same as in Fig. 2. In the left-hand panel, the black hexagons show the median relation from the second *Planck* SZ catalogue (Planck Collaboration et al. 2016), with the error bars denoting 68 per cent of the observed clusters. In the right-hand panel, the black pluses, right-facing triangles, left-facing triangles and hexagons show observational data from Vikhlinin et al. (2006), Sun et al. (2009), Eckmiller et al. (2011) and Planck Collaboration et al. (2011), respectively.

the clusters into \log_{10} mass bins of width 0.2 dex and calculated the median value and 1σ percentiles.

The C-EAGLE clusters provide a consistent extension to the groups and clusters from the periodic volumes. The different AGN calibrations produce negligible difference, even at low-mass group scales. The simulations show good agreement with the observed trend producing a tight power-law relation from low-mass groups to rich clusters, with the greatest outliers being unrelaxed clusters. The normalization of the Y_X relation is marginally higher than the observed relation. This is due to the clusters being too gas rich. The Y_{SZ} relation does not suffer from the problem of clusters being too gas rich, as we measure it within a larger aperture, and the gas fraction averaged within a sphere with such a large radius will tend towards the universal baryon fraction. Our results are consistent with previous numerical work, which have reproduced the observed Y -total mass scaling relations independent of hydrodynamical method and subgrid model (e.g. Battaglia et al. 2012; Kay et al. 2012; Le Brun et al. 2014; Pike et al. 2014; Planelles et al. 2014; Yu, Nelson & Nagai 2015; Gupta et al. 2016; Planelles et al. 2017). For the Y_X relation, the C-EAGLE sample has a scatter of $\sigma_{\log_{10}} = 0.10$, which is similar to the intrinsic scatter of $\sigma_{\log_{10}} = 0.12$ in the HIFLUGCS sample (Eckmiller et al. 2011). If we select relaxed clusters, we observe a tighter power law with a scatter of $\sigma_{\log_{10}} = 0.04$, which is consistent to the scatter of $\sigma_{\log_{10}} = 0.04$ measured by Arnaud, Pointecouteau & Pratt (2007) for a sample of X-ray-selected clusters. The Y_{SZ} -total mass relation shows similar behaviour, with a scatter of $\sigma_{\log_{10}} = 0.18$ for the full sample reducing to $\sigma_{\log_{10}} = 0.06$ for the relaxed sample.

3.4 Summary

We have compared the global properties of the C-EAGLE clusters against observational data and the groups and clusters from two of the periodic volumes from the original EAGLE project. We found

that the C-EAGLE clusters provided a consistent high-mass extension to the periodic volumes for all global scaling relations. We examined the bias introduced by using estimated masses rather than true masses and demonstrated that the assumption of hydrostatic equilibrium introduces a mass independent bias of $b_{\text{hse}} = 0.16$. The use of X-ray estimated profiles increases the bias slightly, leads to greater scatter in the mass estimate and results in a mild mass dependence, with the bias increasing for more massive clusters. Overall the C-EAGLE clusters show reasonable agreement with the observed global scaling relations, reproducing the stellar mass, spectroscopic temperature, X-ray luminosity and SZ signal as a function of total mass relations, the BH mass–stellar mass relation and the global metallicity–spectroscopic temperature relation. The exception is the gas mass–total mass relation where the observed trend is reproduced, but the normalization is too high and the C-EAGLE clusters lie at the top of the intrinsic scatter of the observational data. This impacts the Y_X -total mass relation and its normalization is slightly too high, but the observed trend is reproduced. The higher than observed gas fractions suggest that the AGN feedback is not efficient enough at ejecting gas from the deeper potentials of the progenitors of massive haloes that form at an earlier epoch. Defining a relaxed sample based on the ratio of the kinetic and thermal energy of the hot gas in the ICM, we found that for many relations the scatter is significantly reduced when only considering relaxed clusters.

4 HOT GAS PROFILES

We now examine the hot gas radial profiles of the C-EAGLE clusters to better characterize its distribution and the impact of the AGN feedback model. To make quantitative comparisons to observational data, we must compare like with like. The C-EAGLE sample has a median mass of $M_{500,\text{spec}} = 1.80 \times 10^{14} M_{\odot}$ and median radius of $r_{500,\text{spec}} = 0.88 \text{ Mpc}$. We compare the profiles from the mock X-ray pipeline to the REXCESS cluster sample (Böhringer

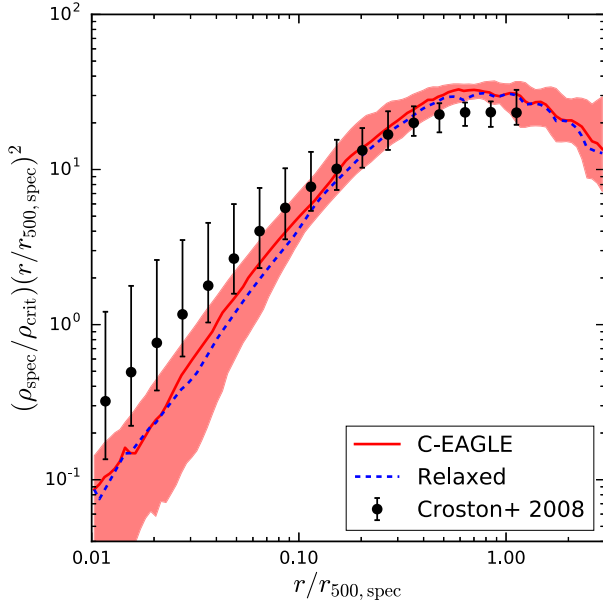


Figure 11. Median density profiles for the C-EAGLE sample (red solid) and the subset of clusters defined as relaxed (blue dashed) at $z = 0.1$, scaled by $(r/r_{500,\text{spec}})^2$ to reduce the dynamic range. The red-shaded region shows the 16th and 84th percentiles of the full sample. The black circles show the median observed profile from the REXCESS sample (Croston et al. 2008), with the error bars enclosing 1σ of the sample.

et al. 2007), but we cut the observational sample such that all clusters have $M_{500} > 10^{14} M_{\odot}$, $z < 0.25$ and the median mass of the sample is $M_{500} = 2.1 \times 10^{14} M_{\odot}$. Matching the median masses of the samples ensures a fairer comparison. However, we do not make any attempt to correct for selection effects. In this section, we calculate dimensionless profiles by dividing by the appropriate quantity, i.e. ρ_{crit} , $k_B T_{500,\text{spec}}$, $P_{500,\text{spec}}$ or $K_{500,\text{spec}}$, where we define these quantities as

$$\rho_{\text{crit}}(z) \equiv E^2(z) \frac{3H_0^2}{8\pi G}, \quad (5)$$

$$k_B T_{500,\text{spec}} = \frac{GM_{500,\text{spec}} \mu m_p}{2r_{500,\text{spec}}}, \quad (6)$$

$$P_{500,\text{spec}} = 500 f_b k_B T_{500,\text{spec}} \frac{\rho_{\text{crit}}}{\mu m_p}, \quad (7)$$

$$K_{500,\text{spec}} = \frac{k_B T_{500,\text{spec}}}{(500 f_b (\rho_{\text{crit}} / \mu_e m_p))^{2/3}}, \quad (8)$$

where H_0 is the Hubble constant and $\mu_e = 1.14$ is the mean atomic weight per free electron.

In Fig. 11, we plot the 3D dimensionless median density profile at $z = 0.1$ for all C-EAGLE clusters and the relaxed subsample. The profiles are scaled by $(r/r_{500,\text{spec}})^2$ to reduce the dynamic range. At $r < 0.1 r_{500,\text{spec}}$ the density is somewhat low compared to the observations, but at radii $> 0.4 r_{500,\text{spec}}$, the hot gas is slightly denser than observed. We also find no significant difference between the relaxed and full samples. The lower than observed density in the cluster cores suggests that the AGN feedback is injecting enough energy to displace and heat central gas and that the injection rate is sufficient to prevent gas from cooling and flowing inwards. However, in Section 3, we found that the C-EAGLE clusters are too

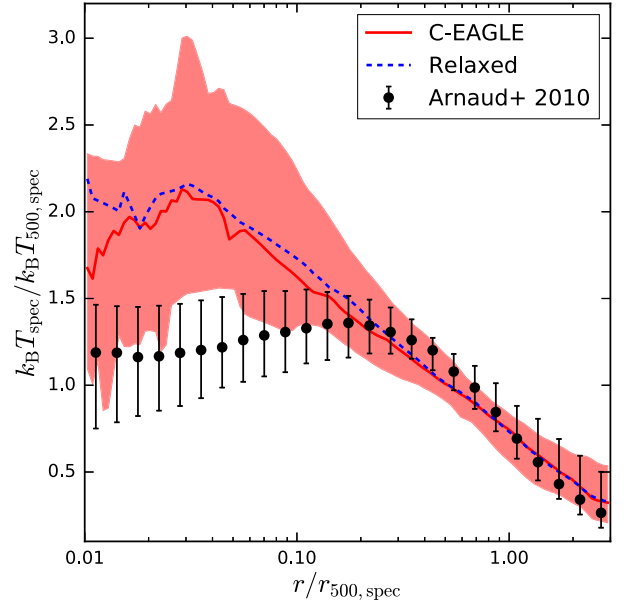


Figure 12. Median spectroscopic temperature profiles for the C-EAGLE sample and the relaxed subset at $z = 0.1$. Marker styles are identical to Fig. 11. The median observed profile is calculated by combining the pressure profiles from Arnaud et al. (2010) with the entropy profiles from Pratt et al. (2010), with the error bars enclosing 1σ of the sample.

gas rich. The excess gas is at radii $> 0.5 r_{500,\text{spec}}$ where the C-EAGLE clusters have a density that exceeds the observed density profile. These results suggest that the AGN feedback in the EAGLE model is capable of ejecting gas from the core of the cluster, but is not moving sufficient amounts beyond r_{500} .

We plot the median dimensionless spectroscopic temperature profile at $z = 0.1$ for the C-EAGLE clusters and the relaxed subset in Fig. 12. Again, we find negligible difference between the full sample and the relaxed sample. At radii greater than $0.2 r_{500,\text{spec}}$, we find good agreement between the observed median profile and the C-EAGLE clusters, with both profiles showing a similar level of scatter. However, inside this radius the two profiles diverge significantly from each other. The observed clusters turn over at $0.2 r_{500,\text{spec}}$, with a roughly flat profile in the cluster core. On the other hand, the C-EAGLE clusters have a median profile that continues to rise until $0.03 r_{500,\text{spec}}$ and reaching a peak value ≈ 60 per cent higher than the observed median profile. The temperature profile is consistent with the density profile, AGN feedback is ejecting and heating central gas in the cluster cores and producing a temperature profile that rises all the way to the centre of the cluster.

In Fig. 13, we plot the median dimensionless pressure profiles, scaled by $(r/r_{500,\text{spec}})^3$, at $z = 0.1$ for the C-EAGLE clusters and the relaxed subset. We find reasonable agreement between the pressure profile of the simulated clusters and the observed pressure profile, with the median profile within the intrinsic scatter of the observed sample at all radii. In the core of the cluster, the profiles begin to diverge, likely due to the injection of energy by the AGN once the cluster core has formed. The agreement of the pressure profiles with the observations, despite clear differences in both the density and temperature profiles, shows that the clusters are approximately in pressure equilibrium. This suggests that central temperature profile is being driven away from the observed profile by constant heating from AGN feedback, rather than capturing the clusters after one large feedback event.

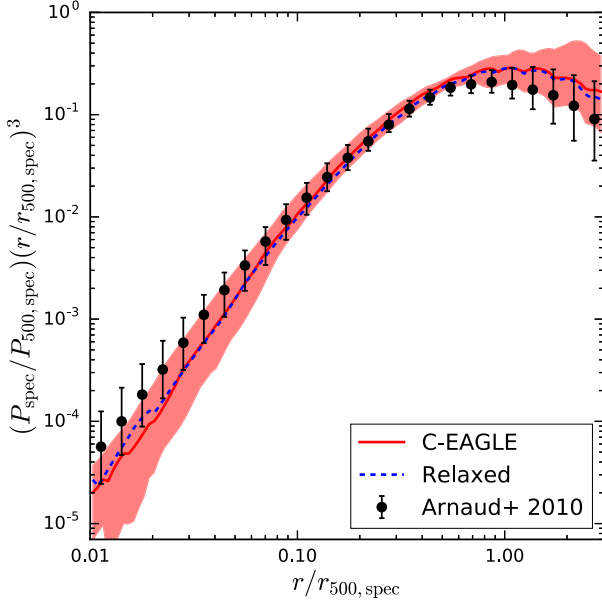


Figure 13. Median pressure profiles for the C-EAGLE sample and the relaxed subset at $z = 0.1$. Marker styles are identical to Fig. 11. Data points show the median observed profile from Arnaud et al. (2010), with the error bars enclosing 1σ of the sample.

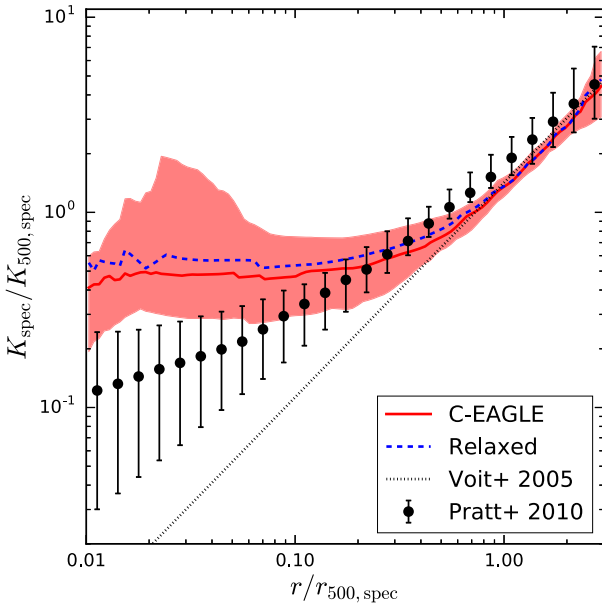


Figure 14. Median entropy profiles for the C-EAGLE sample and its relaxed subset at $z = 0.1$. Marker styles are identical to Fig. 11. Data points show the median observed profile from Pratt et al. (2010), with the error bars enclosing 1σ of the sample. Additionally, we show the prediction from non-radiative simulations (Voit et al. 2005).

We plot the median $z = 0.1$ entropy profiles of the C-EAGLE clusters and the relaxed subset in Fig. 14. Beyond a radius of $0.5 r_{500,\text{spec}}$, the C-EAGLE clusters tend to the power-law result predicted by non-radiative simulations (Voit, Kay & Bryan 2005). However, inside this radius the entropy profile develops a large core. The radial extent of the core region is significantly greater than observed profile and central value of the entropy is up to a factor of 5 larger. Observational results have shown that the central cooling times of clusters take a range of values between those that are significantly

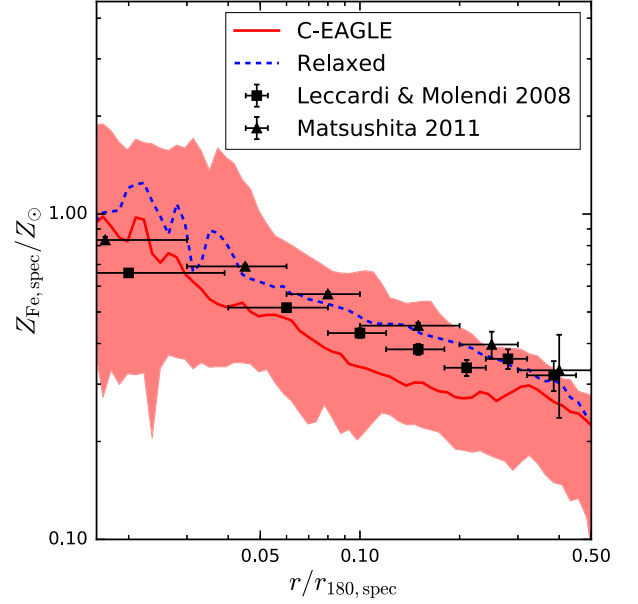


Figure 15. Median emission-weighted iron abundance profiles for the C-EAGLE sample and its relaxed subset at $z = 0.1$. Marker styles are identical to Fig. 11. The black squares and triangles are the mean observed profiles from Leccardi & Molendi (2008) and Matsushita (2011), with the error bars showing the measurement uncertainty.

shorter than the Hubble time, ‘cool-core’ clusters with power-law-like entropy profiles, and those that are longer than the Hubble time, ‘non-cool-core’ clusters with cored entropy profiles (e.g. White, Jones & Forman 1997; Peres et al. 1998; Sanderson, O’Sullivan & Ponman 2009; Cavagnolo et al. 2009). Previous simulation work has reproduced this range of cooling times (Rasia et al. 2015; Hahn et al. 2017). All of the C-EAGLE clusters are non-cool-core clusters with cored entropy profiles. In agreement with the temperature profiles, the entropy profiles suggests that the AGN feedback is heating the core of the cluster. Several of the clusters have inverted central entropy profiles, which can only be produced by a recent injection of energy as it is thermodynamically unstable. This inject of energy by the AGN quickly destroys any cool-core that begins to form.

Finally, we examine the distribution of metals throughout the cluster volume. In Fig. 15, we plot the median emission-weighted iron abundance profiles at $z = 0.1$ for the C-EAGLE clusters and the relaxed subset. We compare with the observations of Leccardi & Molendi (2008) who studied 48 clusters at $z < 0.3$ and Matsushita (2011) who studied 28 clusters at $z < 0.08$. We have scaled all results to the solar abundances of Asplund et al. (2009). For consistency with the observed profiles, we follow Arnaud, Pointecouteau & Pratt (2005) and calculate $r_{180,\text{spec}}$ via

$$r_{180,\text{spec}} = 1.78 \left(\frac{k_B T_{500,\text{spec}}^{\text{X,ce}}}{5 \text{ keV}} \right)^{0.5} E^{-1}(z), \quad (9)$$

where $k_B T_{500,\text{spec}}^{\text{X,ce}}$ is the core-excised spectroscopic X-ray temperature measured by a single temperature fit to the combined spectrum of all particles that fall within $(0.15-1.0) r_{500,\text{spec}}$.

We find reasonable agreement between the observed average metallicity profiles and the median profile of the C-EAGLE sample, though the observations appear to have a marginally flatter profile. The simulated profiles show a similar level of scatter to the observed scatter (see fig. 2 of Leccardi & Molendi 2008). The relaxed subset shows a similar trend to the full sample, but with a normalization

that is systematically higher by 0.1 dex; however, the difference is within the intrinsic scatter of the simulated profiles.

5 DISCUSSION

We have presented the global properties of the C-EAGLE clusters, a set of zoom simulations of 30 objects run with the EAGLE galaxy formation model that yields realistic galaxies and resolves the ICM on kiloparsec scales. We did not recalibrate the model on cluster scales; instead we chose to retain EAGLE’s good match to the observed galaxy stellar mass function, the galaxy size–mass relation and the BH mass–stellar mass relation, thus extending the EAGLE results to higher mass haloes. We found a reasonable match to observations for the total stellar content of the clusters and BH mass–stellar mass relation. The ICM properties, a prediction of a calibrated galaxy formation model, are a reasonable match to the properties of observed low-redshift clusters. The main exceptions are the gas fractions and the cluster entropy profiles. Both of these results suggest that the AGN feedback needs revision if it is to reproduce more realistic cluster properties.

The normalization of the gas fraction–total mass relation is high compared to the observed relation (see Figs 5 and 6). Previous work by McCarthy et al. (2011) has shown that the majority of gas expulsion from galaxy groups and clusters occurs between $z \approx 2$ and 4. In our model, the first AGN feedback events occur when the BH enters its rapid growth phase, which occurs when the host halo becomes massive enough that its halo gas is too hot for supernova feedback to regulate the inflow (Bower et al. 2017). Since the progenitors of higher mass haloes collapse at an earlier epoch than those of smaller haloes, they cross the mass threshold for rapid BH growth earlier when the density of the Universe is higher. This means that their potential wells are deeper, requiring more energetic feedback to expel the same fraction of gas. The higher-than-observed gas fractions of the C-EAGLE clusters suggest that AGN feedback does not inject sufficient energy into cluster progenitors at high redshift. This may also explain why the BCGs are too massive (Bahé et al. 2017), since more efficient feedback at higher redshift will also suppress star formation in cluster progenitors.

In addition, a comparison of the simulated and observed entropy profiles (Fig. 14) shows that the simulated clusters have significantly larger cores, higher central entropies than observed and none of them are cool cores. This suggests that AGN feedback continues injecting energy into the cluster cores once they have formed, maintaining or increasing their size and preventing cool cores from ever reforming. AGN feedback in EAGLE uses a modified Booth & Schaye (2009) model, which is based on Springel et al. (2005). These two models have been used by several previous studies (Planelles et al. 2014; Pike et al. 2014; Le Brun et al. 2014; McCarthy et al. 2017) that broadly reproduce the observed entropy profiles of clusters. The principal difference between this study and previous work is that we achieve a mass resolution that is a factor of 100–500 times better. This lowers the energy threshold for an AGN feedback event by the same factor in our implementation. The decrease in energy threshold results in individual events that are less energetic but occur more frequently, thus rendering the feedback smoother and less episodic. In addition, the high resolution of the C-EAGLE simulations means that the gas reaches higher densities around the central BH compared to previous studies. In a Bondi–Hoyle model, the rate of accretion scales as the mass of the BH squared. Therefore, once the cluster core has formed and the central BH is massive, even modest larger scale densities will be enhanced in the immediate vicinity of the BH and so lead to significant accretion. Thus, it is not necessary for a

cool core to reform in order to trigger an AGN feedback event, as is the case at lower resolution.

If the AGN feedback events were more energetic, they might expel gas more effectively from high-redshift progenitors, potentially making gas infall and star formation less efficient. Increasing the energy threshold for individual events would also make feedback more intermittent, perhaps allowing cool cores to reform during periods of low activity. However, in the current EAGLE model we are limited by the pressure–entropy SPH formalism, which fails to conserve energy if too many particles are heated in a single time step. In addition, there are other effects to consider. The AGN feedback model currently distributes heat isotropically to the surrounding gas. If the heating was concentrated into a jet, it might allow a cool core to coexist with the transport of energy to larger radii as is observed, for example, in the Perseus cluster (e.g. Zhuravleva et al. 2014). The role of our artificial conduction scheme in shaping the cluster core is also unclear. Improvement of the EAGLE AGN feedback model is ongoing, but this work highlights that we must continue to improve the modelling of processes that occur below the resolution limit of the simulations as we push to higher resolution.

6 SUMMARY

In this paper, we have introduced the Cluster EAGLE (C-EAGLE) simulation project, a set of zoom simulations of the formation of 30 galaxy clusters in the mass range of $10^{14} < M_{200}/M_{\odot} < 10^{15.4}$ that resolve the cluster galaxies and the ICM on kiloparsec scales using the state-of-the-art EAGLE AGNdt9 model of S15. In this paper, we have presented the global cluster properties, a prediction of a model calibrated for the formation of field galaxies. The properties of the cluster galaxies are analysed in Bahé et al. (2017). Our main results are as follows:

- (i) The C-EAGLE clusters provide a consistent extension to periodic volumes of the original EAGLE project, sampling the massive objects that were missing due to their limited volumes.
- (ii) Estimating masses from mock X-ray observations rather than using the true mass results in a hydrostatic bias of 15–20 per cent (Fig. 2). Selecting relaxed clusters produces a similar bias but significantly reduces the scatter.
- (iii) The clusters show reasonable agreement with the observed relations that were used to calibrate the EAGLE model at lower masses. The clusters show a flat stellar fraction–total mass relation in agreement with the observations (Fig. 3), and the trend, normalization and scatter of the BH mass–stellar mass relation are also in agreement with the observations (Fig. 4). The C-EAGLE clusters contain a population of orphan supermassive BHs, likely the remnants of galaxy mergers. Several of these orphans reside inside the BCGs.
- (iv) The C-EAGLE clusters reproduce the observed gas fraction–total mass trend, but the normalization is too high by ≈ 30 per cent (Fig. 5). Although we selected the AGNdt9 calibration as it better reproduced the gas fractions of low-mass galaxy groups, the different EAGLE AGN calibrations produce similar results on cluster scales. The progenitors of clusters will form at an earlier epoch than those of less massive haloes, causing them to have deeper potentials due to the increased density of the Universe. The increased gas fractions suggests that in this regime the EAGLE AGN feedback model does not inject sufficient energy into the progenitors of the C-EAGLE clusters.
- (v) The predicted global X-ray and SZ properties of the clusters (Figs 7–10) show a good match to the spectroscopic temperature, X-ray luminosity and Y_{SZ} total mass relations. They reproduce the

trend of the Y_X –mass relation, but the normalization is marginally high due to the high gas fractions.

(vi) The C-EAGLE clusters show good agreement with observed metallicities (Fig. 9) and their distribution throughout the cluster volume (Fig. 15).

(vii) Median hot gas profiles show that our cluster cores are uniformly less dense and hotter than observed (Figs 11–14). As a result, our clusters have larger than observed entropy cores and our sample includes no cool core clusters. This suggests that the AGN feedback is insufficiently episodic, injecting energy continually into hot gas cores once they have formed. The median pressure profile is in reasonable agreement with the observed profile, showing that the differences between the simulations and the observations are primarily in the thermodynamic rather than in the dynamical state of the gas.

The C-EAGLE clusters are an attempt to simulate the formation of rich galaxy clusters with a model that yields a galaxy population that is a good match to the observed field population. The resolution required to produce realistic galaxies is more than an order of magnitude higher than that in most previous cluster studies. Overall, the global properties of the ICM, such as the temperature, luminosity, and the total content and distribution of metals, are in reasonable agreement with the observations. A companion paper (Bahé et al. 2017) demonstrates that the galaxy properties are also in reasonable agreement with observations, although our BCGs are clearly too massive. Future C-EAGLE papers will address the formation and evolution of cluster galaxies, the interaction of galaxies with the ICM, the formation of structures in the ICM, and mass estimation systematics for cluster cosmology. We will also investigate the effects of the AGN feedback in more detail, with the aim of producing an improved model that reproduces even more closely the observed properties of cluster galaxies and the ICM.

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APPENDIX A: TABLE OF OBSERVABLE QUANTITIES

The C-EAGLE sample was selected from a large dark matter only simulation and resimulated using the zoom technique. To enable a fair comparison of the simulations with observational data, we compute a mock X-ray spectrum for every particle in the simulation. These are then summed to provide a mock spectrum for the cluster. We

Table A1. Global quantities of the C-EAGLE clusters at $z = 0.1$, with ‘hse’ and ‘spec’ quantities estimated via the mock X-ray pipeline. The IC location and IC extent specify the comoving centre and radius of the sphere that contains the high-resolution Lagrangian region in the initial conditions. The PANPHASIA phase descriptor is [Panh1, L14, (2152, 5744, 757), S3, CH1814785143, EAGLE_L3200_VOL1].

Cluster	$M_{500,\text{true}}$	$M_{500,\text{hse}}$ ($\log_{10}(M/M_{\odot})$)	$M_{500,\text{spec}}$	$r_{500,\text{true}}$	$r_{500,\text{spec}}$ (Mpc)	x (Mpc h^{-1})	IC location			IC extent [Mpc h^{-1}]
							y (Mpc h^{-1})	z (Mpc h^{-1})		
CE-00	13.905	13.756	13.751	0.65	0.58	207.81	1498.48	1793.66		22.45
CE-01	13.982	13.927	13.931	0.69	0.66	1765.99	1721.52	1541.50		21.03
CE-02	13.920	13.870	13.871	0.66	0.63	1962.55	1953.81	234.28		21.03
CE-03	13.926	13.877	13.861	0.66	0.63	1772.74	1915.43	616.81		27.33
CE-04	13.853	13.723	13.697	0.62	0.55	1170.49	1524.83	1807.16		23.97
CE-05	13.892	13.902	13.880	0.64	0.64	395.90	613.11	1126.81		23.97
CE-06	14.128	14.102	14.083	0.77	0.74	1774.60	1519.13	206.88		35.53
CE-07	14.176	14.127	14.121	0.80	0.77	856.66	1662.35	869.10		31.16
CE-08	14.070	13.980	13.941	0.74	0.67	329.85	497.42	247.39		24.63
CE-09	14.244	14.118	14.084	0.84	0.74	922.84	982.93	1499.01		24.77
CE-10	14.301	14.250	14.225	0.88	0.83	1771.02	1093.98	1271.48		25.60
CE-11	14.290	14.182	14.182	0.87	0.80	1740.71	459.04	920.32		26.45
CE-12	14.422	14.349	14.326	0.96	0.90	793.25	945.74	682.99		36.72
CE-13	14.341	14.249	14.266	0.91	0.86	682.60	1023.43	1327.07		32.20
CE-14	14.563	14.571	14.568	1.08	1.08	184.91	991.14	1381.74		49.32
CE-15	14.407	14.404	14.418	0.95	0.96	1364.14	501.66	1179.09		41.86
CE-16	14.311	14.170	14.143	0.89	0.78	487.63	1526.81	406.49		41.86
CE-17	14.537	14.461	14.433	1.05	0.97	143.61	1251.36	1953.15		31.16
CE-18	14.639	14.587	14.555	1.14	1.07	542.69	580.55	1091.34		41.86
CE-19	14.586	14.483	14.445	1.09	0.98	547.19	219.86	760.69		37.94
CE-20	14.482	14.361	14.342	1.01	0.91	1827.41	1204.77	2009.96		30.16
CE-21	14.800	15.012	14.934	1.29	1.43	767.31	607.95	650.43		40.51
CE-22	14.837	14.721	14.701	1.33	1.20	1407.02	1572.74	567.84		62.05
CE-23	14.426	14.291	14.256	0.97	0.85	1378.96	2033.10	1838.00		37.94
CE-24	14.821	14.728	14.666	1.31	1.16	209.13	668.70	1948.65		52.66
CE-25	15.045	15.070	15.068	1.56	1.58	697.69	861.16	860.76		49.32
CE-26	14.899	14.838	14.780	1.39	1.27	1907.88	852.55	1346.40		43.26
CE-27	14.689	14.496	14.389	1.18	0.94	1790.48	618.80	1783.99		41.86
CE-28	14.902	14.688	14.671	1.39	1.17	944.94	707.08	1388.09		62.05
CE-29	15.077	15.089	14.912	1.60	1.40	719.79	1449.38	1015.75		60.04

Table A2. Global quantities of the C-EAGLE clusters at $z = 0.1$, estimated from the mock X-ray pipeline. All quantities are measured inside $r_{500,\text{spec}}$ except for Y_{SZ} , which is measured inside $5r_{500,\text{spec}}$.

Cluster	$k_{\text{B}}T^{\text{X}}$ (keV)	$L_{\text{X}}^{0.5-2.0\text{keV}}$ ($\log_{10}(L/\text{erg s}^{-1})$)	M_{gas} ($\log_{10}(M/M_{\odot})$)	M_{star} ($\log_{10}(M/M_{\odot})$)	Y_{X} ($\log_{10}(Y/M_{\odot}\text{keV})$)	Y_{SZ} ($\log_{10}(Y/\text{Mpc}^2)$)	Z_{Fe} ($Z_{\text{Fe},\odot}$)	$E_{\text{kin}}/E_{\text{thrm}}$
CE-00	1.81	43.234	12.881	12.230	13.125	−5.445	0.26	0.11
CE-01	2.18	43.459	13.026	12.290	13.302	−5.521	0.21	0.24
CE-02	1.90	43.227	12.954	12.164	13.217	−5.425	0.27	0.05
CE-03	2.01	43.114	12.924	12.140	13.221	−5.298	0.21	0.06
CE-04	1.79	42.888	12.789	11.994	13.027	−5.449	0.22	0.11
CE-05	1.75	43.240	12.970	12.154	13.213	−5.340	0.29	0.09
CE-06	2.28	43.518	13.204	12.306	13.551	−5.046	0.22	0.09
CE-07	2.48	43.657	13.247	12.354	13.631	−4.984	0.21	0.09
CE-08	2.03	43.518	13.120	12.247	13.412	−5.126	0.25	0.09
CE-09	2.60	43.620	13.237	12.434	13.641	−4.850	0.20	0.16
CE-10	3.00	43.890	13.414	12.479	13.886	−4.758	0.26	0.05
CE-11	2.66	44.080	13.389	12.482	13.815	−4.845	0.29	0.21
CE-12	3.28	43.942	13.485	12.644	13.990	−4.636	0.26	0.08
CE-13	3.20	44.117	13.443	12.460	13.917	−4.721	0.29	0.06
CE-14	3.99	44.321	13.704	12.709	14.302	−4.508	0.23	0.31
CE-15	3.56	44.037	13.511	12.592	13.993	−4.608	0.16	0.22
CE-16	3.06	43.877	13.351	12.505	13.814	−4.522	0.26	0.12
CE-17	3.91	44.284	13.623	12.684	14.208	−4.525	0.21	0.25
CE-18	4.22	44.402	13.744	12.787	14.359	−4.262	0.19	0.09
CE-19	3.23	44.102	13.660	12.723	14.171	−4.354	0.25	0.30
CE-20	3.60	44.127	13.551	12.660	14.077	−4.457	0.24	0.12

Table A2 – *continued*

Cluster	$k_B T^X$ (keV)	$L_X^{0.5-2.0\text{keV}}$ ($\log_{10}(L/\text{erg s}^{-1})$)	M_{gas} ($\log_{10}(M/M_{\odot})$)	M_{star} ($\log_{10}(M/M_{\odot})$)	Y_X ($\log_{10}(Y/M_{\odot}\text{keV})$)	Y_{SZ} ($\log_{10}(Y/\text{Mpc}^2)$)	Z_{Fe} ($Z_{\text{Fe},\odot}$)	$E_{\text{kin}}/E_{\text{thrm}}$
CE-21	5.09	44.540	13.990	13.024	14.694	−4.009	0.51	0.29
CE-22	6.04	44.618	13.937	12.967	14.668	−3.854	0.16	0.12
CE-23	3.16	43.904	13.452	12.642	13.934	−4.325	0.19	0.26
CE-24	5.31	44.412	13.876	12.937	14.598	−4.005	0.17	0.13
CE-25	7.95	45.021	14.214	13.177	15.102	−3.650	0.23	0.31
CE-26	6.43	44.589	14.010	13.037	14.809	−3.765	0.19	0.08
CE-27	5.28	42.987	13.094	12.323	13.811	−4.259	0.19	0.24
CE-28	6.29	44.407	13.906	13.051	14.697	−3.735	0.18	0.13
CE-29	7.66	44.942	14.188	13.185	15.067	−3.510	0.52	0.30

then perform a hydrostatic analysis using the density and temperature profiles measured from the mock spectra to produced estimated masses and radii. Observable quantities are then calculated within estimated apertures, see Le Brun et al. (2014) for further details. The tables below provide the locations of the sample in the parent simulation $z = 0$, the size of the sphere selected at $z = 0$ from the parent to produce the high-resolution region and the global properties of the clusters at $z = 0.1$ that are presented in Section 3. For additional cluster properties, we direct the reader to Appendix A of Bahé et al. (2017).

APPENDIX B: POWER-LAW FITS

In this work, we have presented how observable quantities scale with estimated mass. To get a quantitative measure of the scatter of the C-EAGLE clusters about each scaling relation, we assume that the cluster scaling relations are simple power laws (Kaiser 1986). At $z = 0.1$, we fit each scaling relation with a model of the

form

$$Y = 10^A \left(\frac{X}{X_{\text{piv}}} \right)^\alpha, \quad (\text{B1})$$

where Y is the observable quantity, X is the total mass, A and α are the normalization and slope of the power law and $X_{\text{piv}} = 4.0 \times 10^{14} M_{\odot}$ is the pivot mass. The scatter about the power-law relation is then calculated via

$$\sigma_{\log_{10}} = \sqrt{\frac{1}{N} \sum_{i=1}^N [\log_{10}(Y_i) - \log_{10}(Y_{\text{mod}})]^2}, \quad (\text{B2})$$

where $N = 30$ is the number of clusters in the sample, Y_i is the measured observable quantity, Y_{mod} is the expected observable value of a cluster of mass X_i and $\sigma_{\ln} = \ln(10)\sigma_{\log_{10}}$. We fit power laws to the full sample, those clusters defined as relaxed and those defined as unrelaxed. Table B1 summarizes the power-law best-fitting relations and the scatter about them.

Table B1. Summary of the best-fitting values of the linear power-law fits and the scatter about each fit for the scaling relations presented in this work. All fits are done at $z = 0.1$. The values of A given in the table assume that the units of 10^A are keV, ergs^{-1} , M_{\odot} , M_{\odot} , M_{\odot} keV and Mpc^2 , respectively.

Scaling relation	Full sample			Relaxed			Unrelaxed		
	A	α	$\sigma_{\log_{10}}$	A	α	$\sigma_{\log_{10}}$	A	α	$\sigma_{\log_{10}}$
$k_B T_{500,\text{spec}}^X - M_{500,\text{spec}}$	$0.67^{+0.03}_{-0.02}$	$0.47^{+0.07}_{-0.02}$	0.06	$0.68^{+0.01}_{-0.02}$	$0.56^{+0.02}_{-0.03}$	0.02	$0.68^{+0.02}_{-0.02}$	$0.42^{+0.09}_{-0.02}$	0.07
$L_{X,500,\text{spec}}^{0.5-2.0\text{keV}} - M_{500,\text{spec}}$	$44.13^{+0.24}_{-0.07}$	$1.33^{+0.13}_{-0.08}$	0.30	$44.45^{+0.11}_{-0.06}$	$1.59^{+0.21}_{-0.08}$	0.11	$44.11^{+0.26}_{-0.08}$	$1.36^{+0.12}_{-0.15}$	0.35
$M_{\text{gas},500,\text{spec}} - M_{500,\text{spec}}$	$13.68^{+0.09}_{-0.03}$	$1.07^{+0.02}_{-0.05}$	0.13	$13.82^{+0.01}_{-0.01}$	$1.15^{+0.03}_{-0.02}$	0.03	$13.68^{+0.09}_{-0.03}$	$1.09^{+0.03}_{-0.08}$	0.15
$M_{\text{star},500,\text{spec}} - M_{500,\text{spec}}$	$12.79^{+0.05}_{-0.03}$	$0.85^{+0.06}_{-0.04}$	0.10	$12.85^{+0.02}_{-0.02}$	$0.97^{+0.03}_{-0.03}$	0.03	$12.79^{+0.06}_{-0.03}$	$0.85^{+0.06}_{-0.06}$	0.12
$Y_{X,500,\text{spec}} - M_{500,\text{spec}}$	$14.39^{+0.03}_{-0.02}$	$1.57^{+0.07}_{-0.07}$	0.10	$14.48^{+0.01}_{-0.01}$	$1.69^{+0.04}_{-0.02}$	0.04	$14.39^{+0.04}_{-0.02}$	$1.57^{+0.07}_{-0.10}$	0.12
$Y_{\text{SZ},500,\text{spec}} - M_{500,\text{spec}}$	$-4.16^{+0.04}_{-0.03}$	$1.48^{+0.11}_{-0.10}$	0.18	$-4.13^{+0.01}_{-0.06}$	$1.66^{+0.03}_{-0.09}$	0.06	$-4.16^{+0.05}_{-0.03}$	$1.46^{+0.12}_{-0.17}$	0.21

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